Implications of M_t for ElectroWeak Vacuum Stability

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IFAE

OUTLINE

- Status after first LHC run
  Higgs discovered, no trace of BSM...

- $M_h \simeq 126$ GeV $\Rightarrow$ EW vacuum unstable
  The impact of Δm_t

- Several implications of this instability

REFERENCES

EARLY WORK ON VACUUM INSTABILITY

I. Krive, A. Linde '76

N. Krasnikov '78

L. Maiani, G. Parisi, R. Petronzio '78 + N. Cabibbo '79

H. Politzer, S. Wolfram '79

P. Hung '79

A. Linde '80

M. Lindner '86 + M. Sher, H. Zaglauer '89

+ ... many more

REFERENCES

RECENT PRECISION STUDIES

... +

M. Holthausen, K.S. Lim, M. Lindner [ph/1112.2415]

J. Elias-Miró, JRE, G.F. Giudice, G. Isidori, A. Riotto, A. Strumia [ph/1112.3022]

F. Bezrukov, M.Y. Kalmykov, B.A. Kniehl, M. Shaposhnikov [ph/1205.2893]

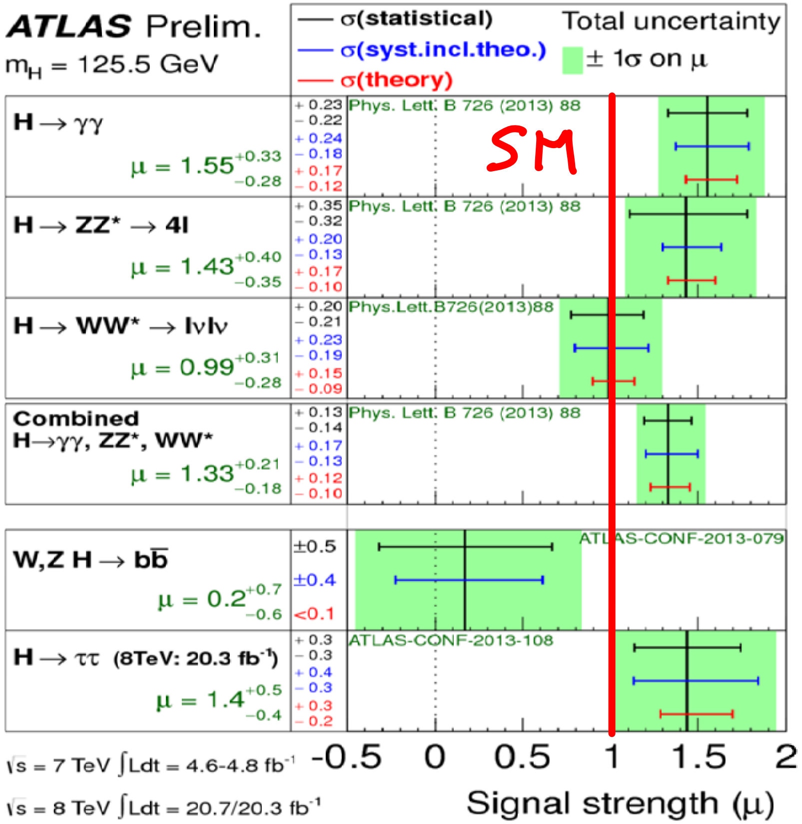
G. Degrassi, S. Di Vita, J. Elias-Miró, JRE, G.F. Giudice, G. Isidori, A. Strumia [ph/1205.6497]

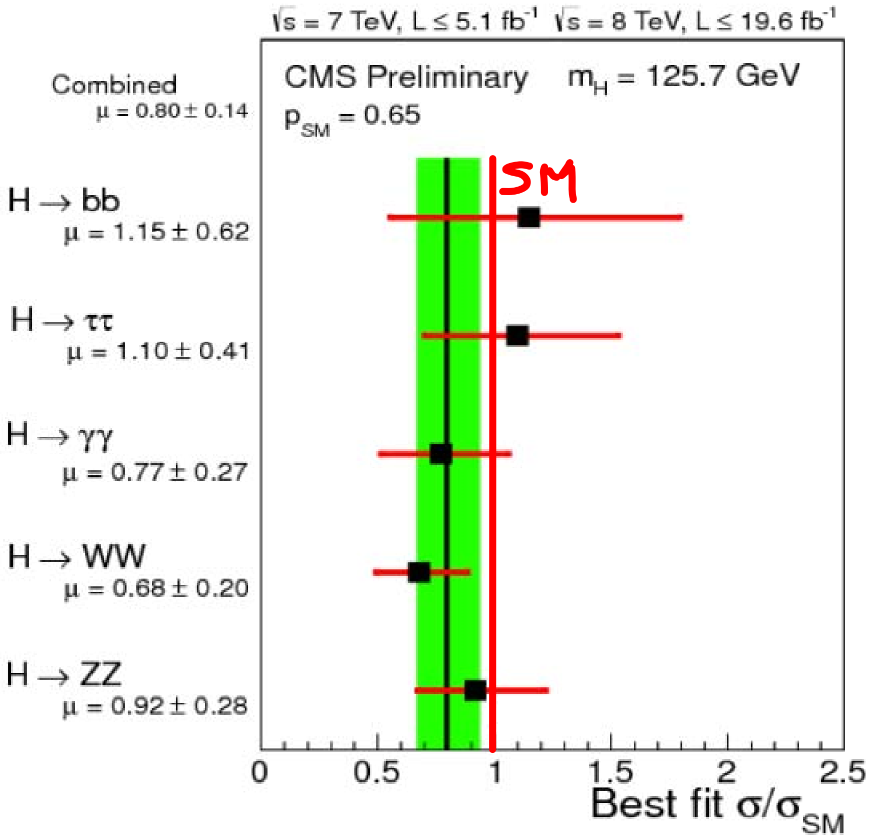
S. Alekhin, A. Djouadi, S. Moch [ph/1207.0980]

D. Buttazzo, G. Degrassi, P. Giardino, G. Giudice, F. Sala, A. Salvio, A. Strumia [ph/1307.3536]

SM STATUS

- Higgs discovered, close to SM-like





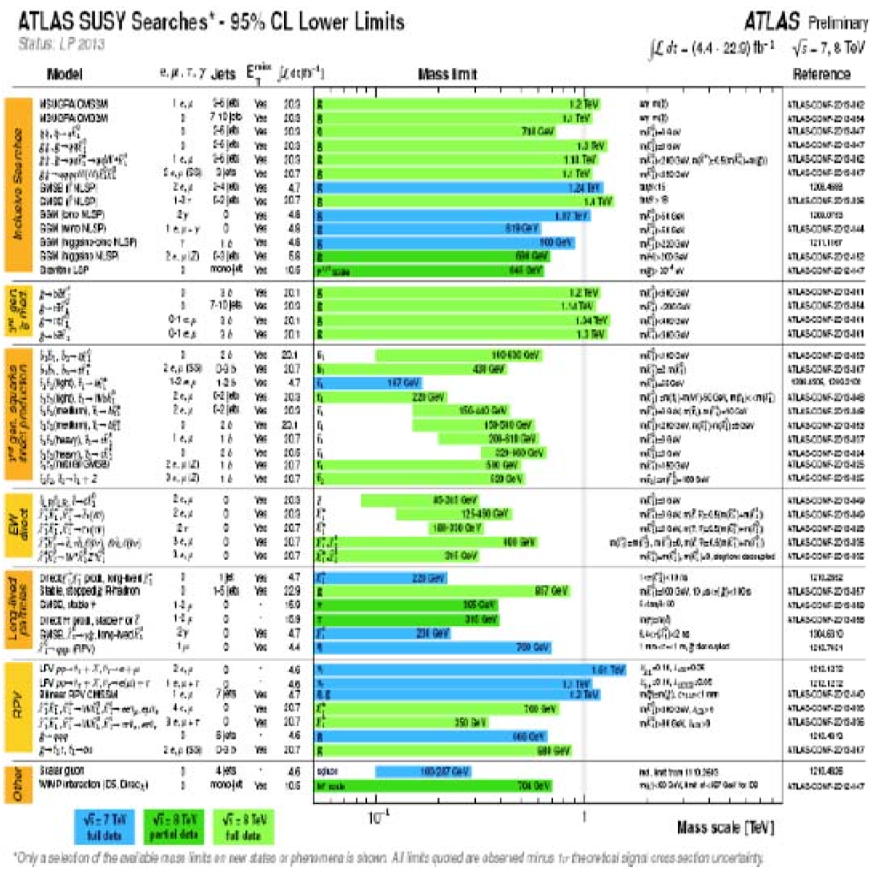
Mh/GeV = 125.5 + 0.2 (stat) + 0.5/-0.6 (syst) ATLAS

Mh/GeV = 125.7 + 0.3 (stat) + 0.3 (syst) CMS

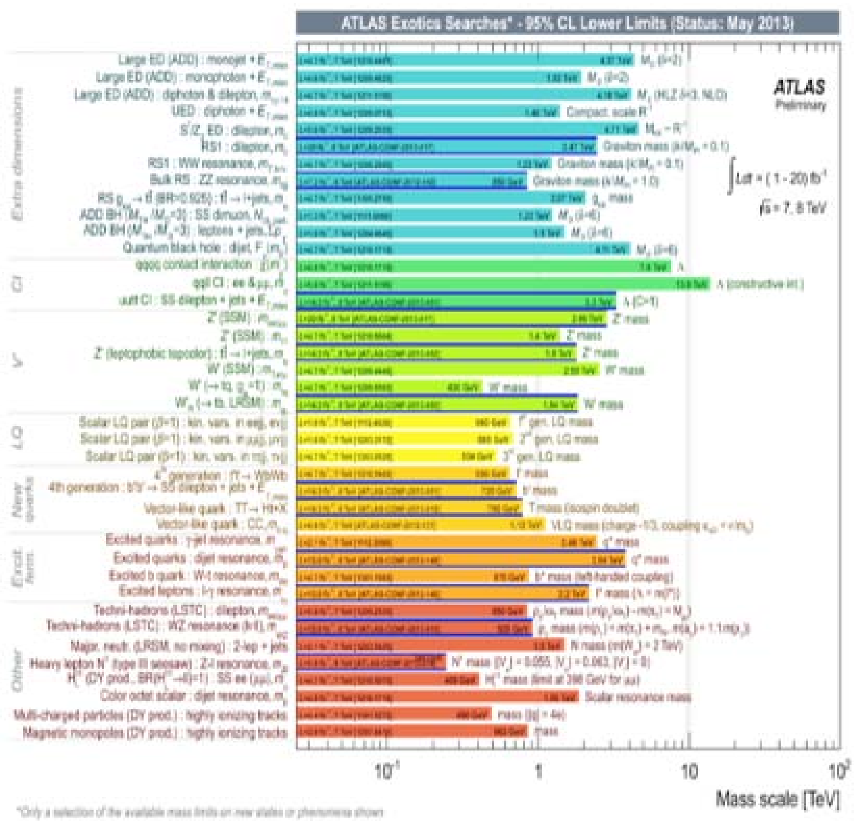
BSM STATUS

- No trace of BSM so far ⇒ $\Lambda >$ few TeV ?

"TSUNAMI" EXCLUSION PLOTS

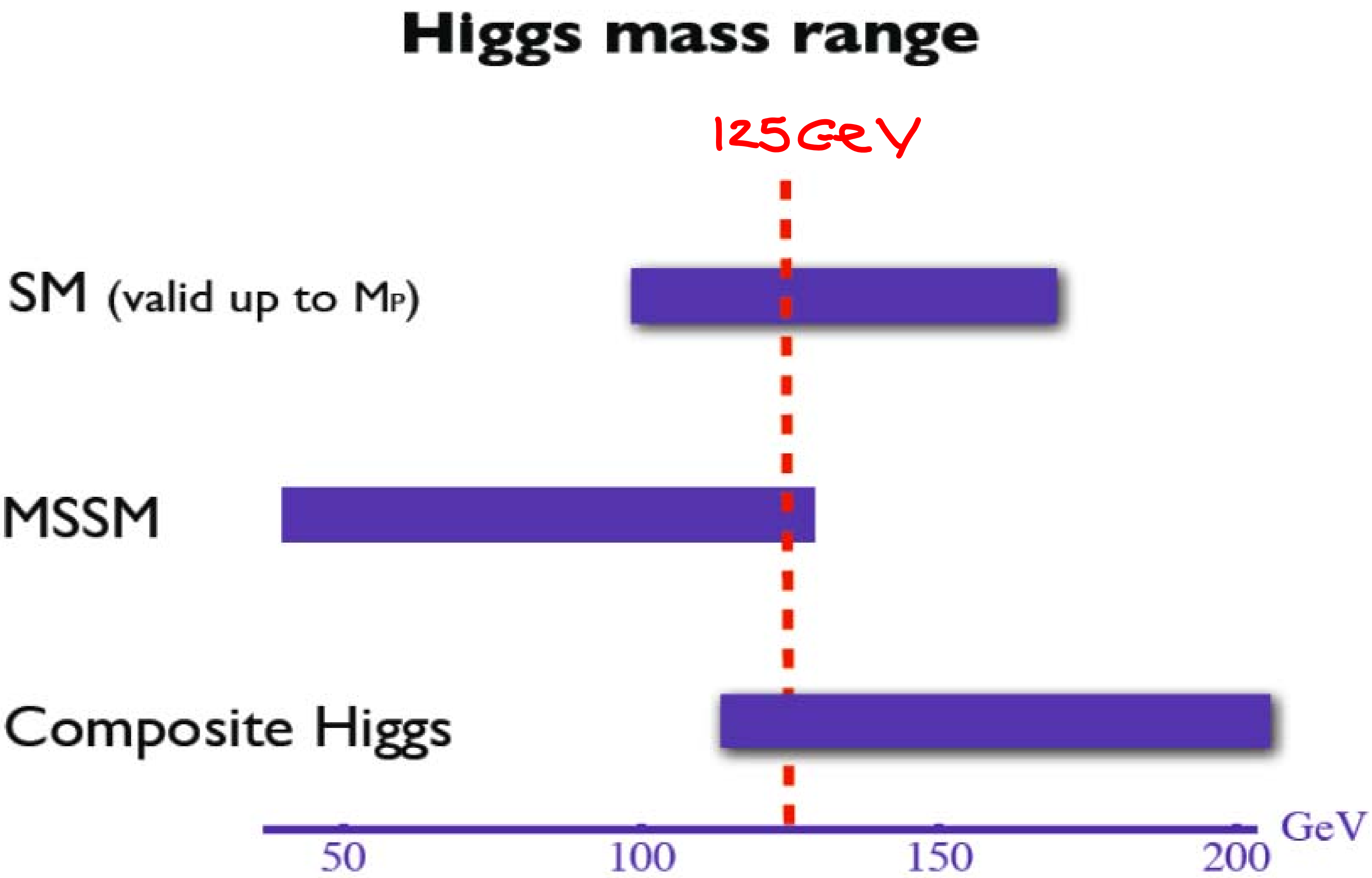


SUSY



EXOTICS

M_h AS MODEL DISCRIMINATOR



A. Pomarol, ICHEP'12

BSM STATUS

- Higgs discovered, close to SM-like

+

- No trace of BSM so far ⇒ $\Lambda >$ few TeV ?

+

- Holding on to naturalness

⇓

$\Lambda \sim$ few TeV

BSM STATUS / THIS TALK

- Higgs discovered, close to SM-like

+

- No trace of BSM so far ⇨ $\Lambda \gg$ few TeV ?

+

- Disregarding naturalness

⇩

$\Lambda \sim M_{Pl}$?

SM EXTRAPOLATION

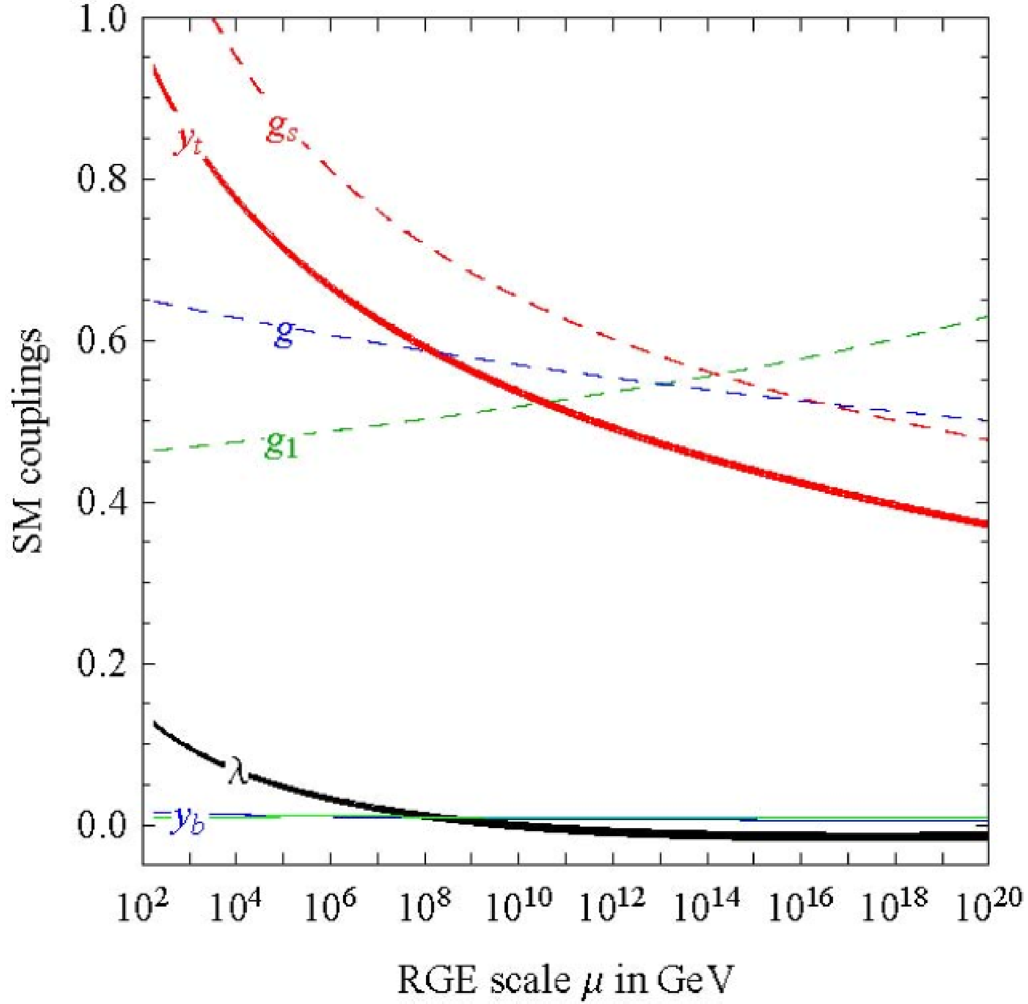
Assume Higgs has SM props. and no BSM Physics

All SM parameters known

$M_h \Rightarrow \lambda(EW)$

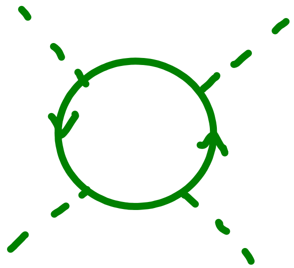
Forgetting naturalness, can the pure SM be valid up to M_{Pl}?

Weakly coupled up to M_{Pl}



VACUUM INSTABILITY

$$\frac{d\lambda}{d\ln Q} \sim -\frac{h_t^4}{16\pi^2}$$

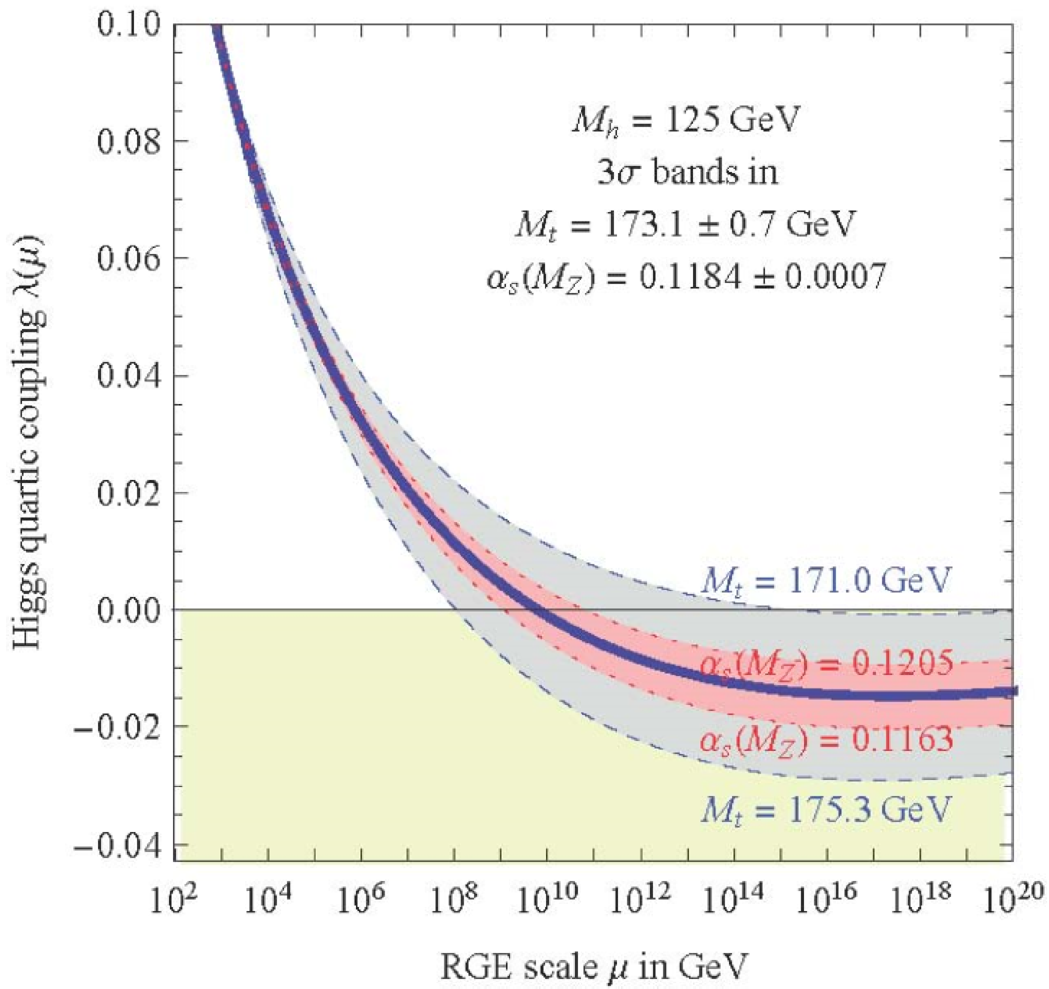


$\lambda < 0$ at $\Lambda_I \sim 10^{10}$ GeV

⇓

Higgs potential instability

$$V(\phi \gg M_t) \simeq \frac{1}{4}\lambda(Q \simeq h)\, h^4$$



Cabibbo et al '79, Hung '79, Lindner '86

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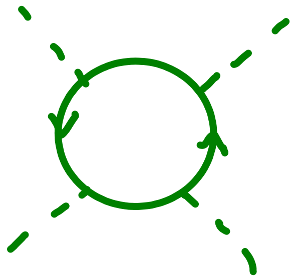

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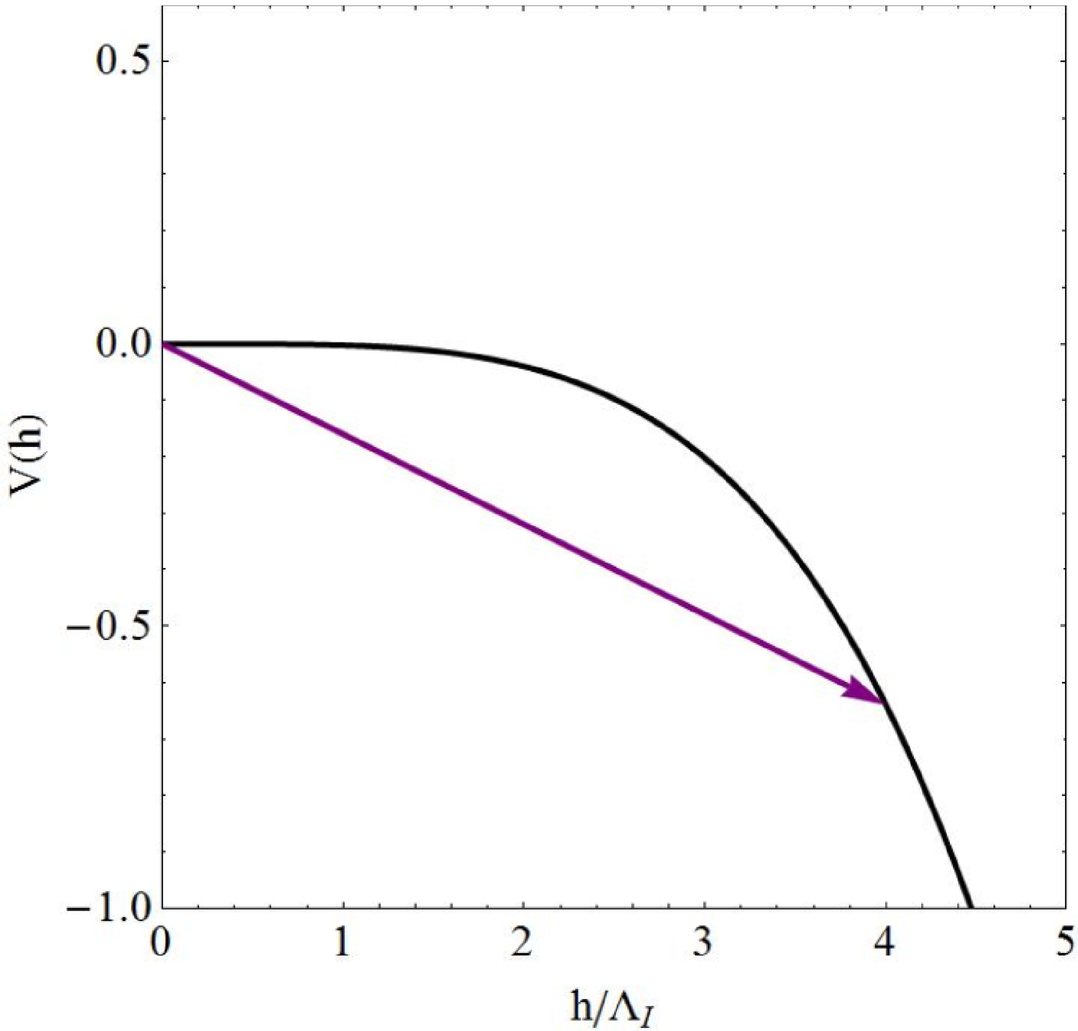

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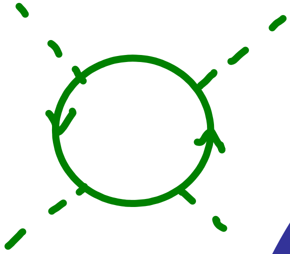

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⬇

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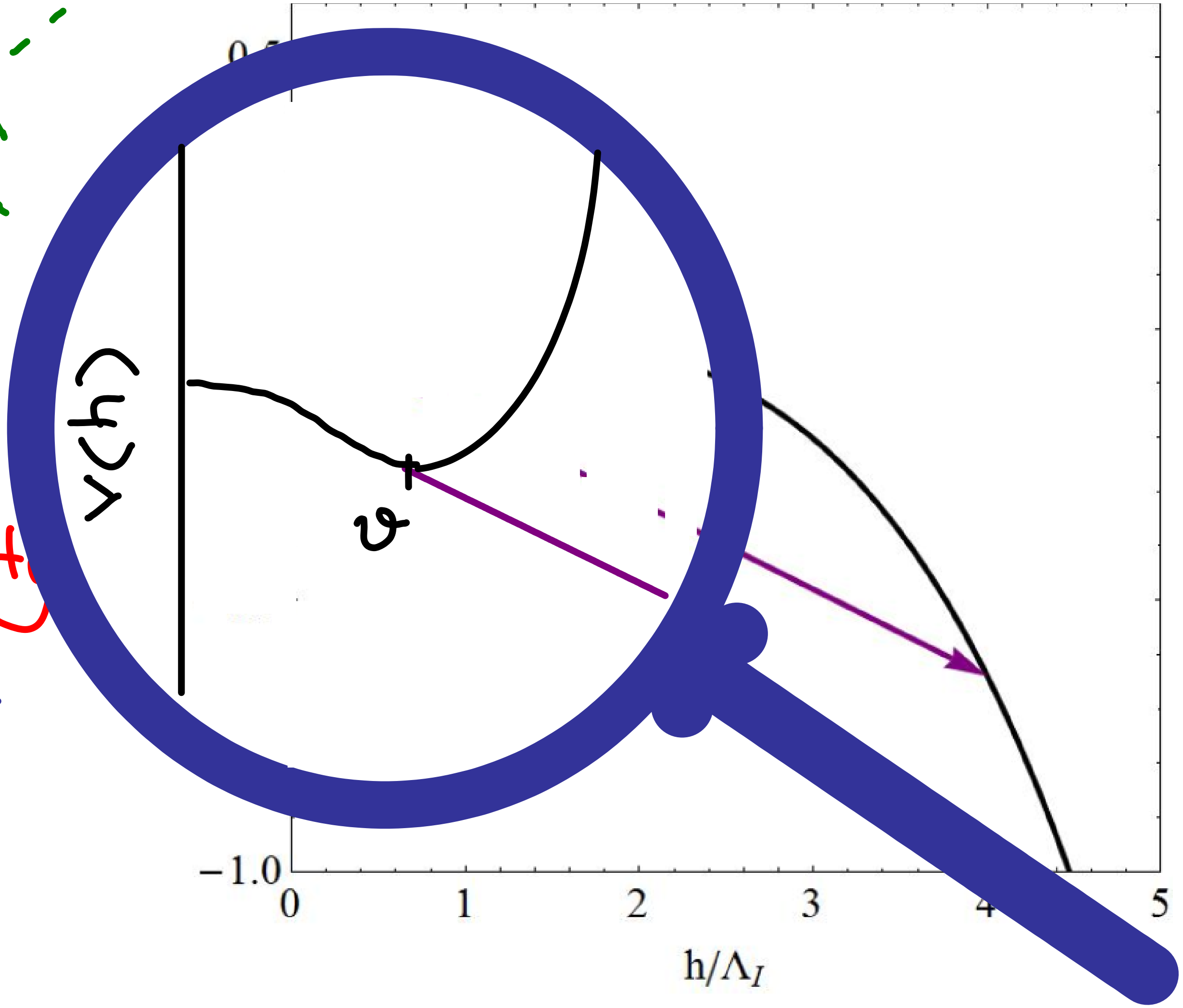


LIFE IN A METASTABLE VACUUM

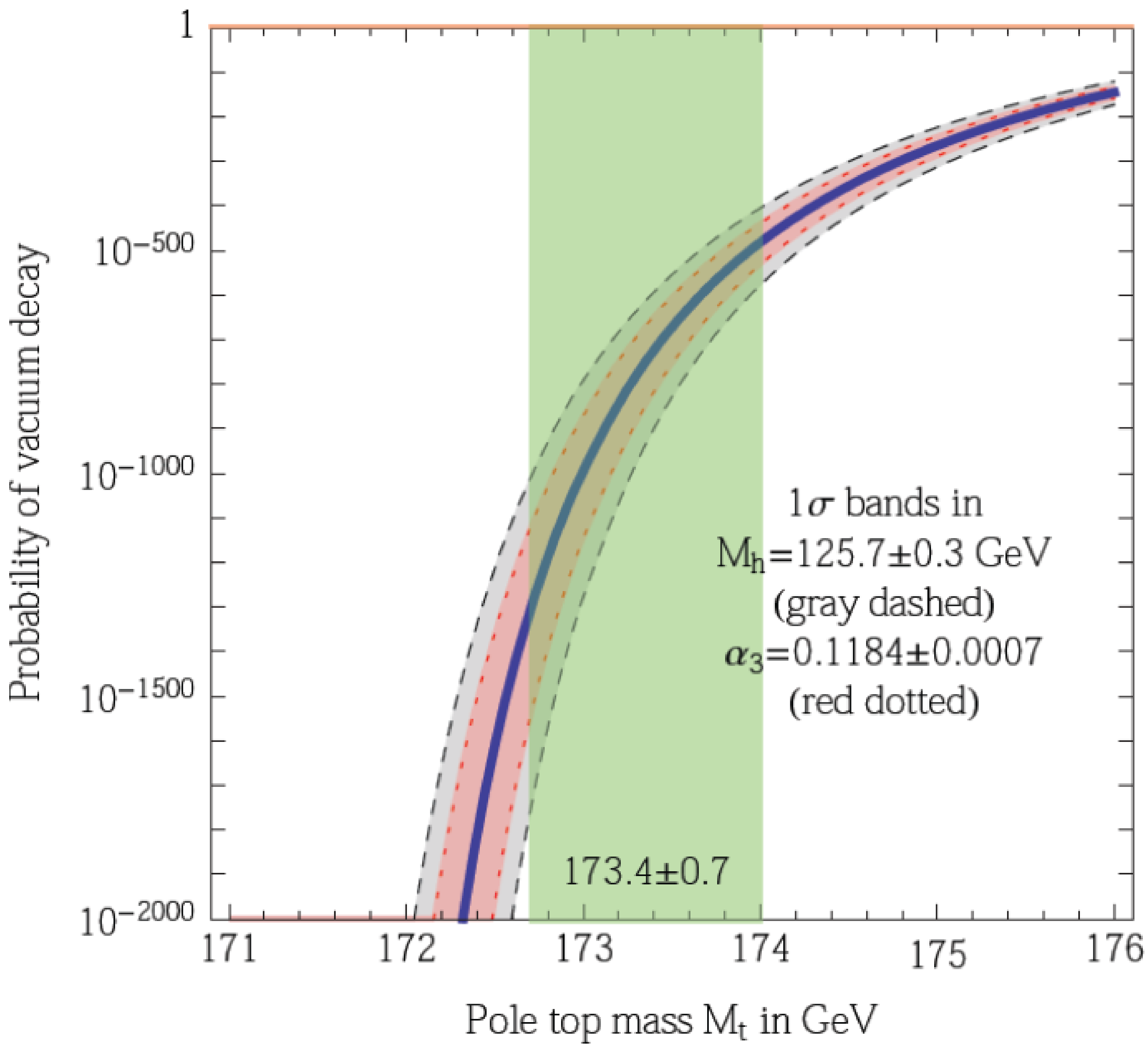
$$p = \text{Decay prob.} = \underbrace{\frac{\text{Decay rate}}{\Delta t \cdot \Delta V}}_{h^4 e^{-S_4}} \tau_U^4 \quad \text{with } \tau_U^4 \sim \left(e^{140}/M_{Pl}\right)^4$$

$$h^4 e^{-S_4} \sim h^4 \exp\left(-\frac{8\pi^2}{3|\lambda(h)|}\right) \sim h^4 \exp\left[-\frac{2600}{|\lambda|/0.01}\right]$$

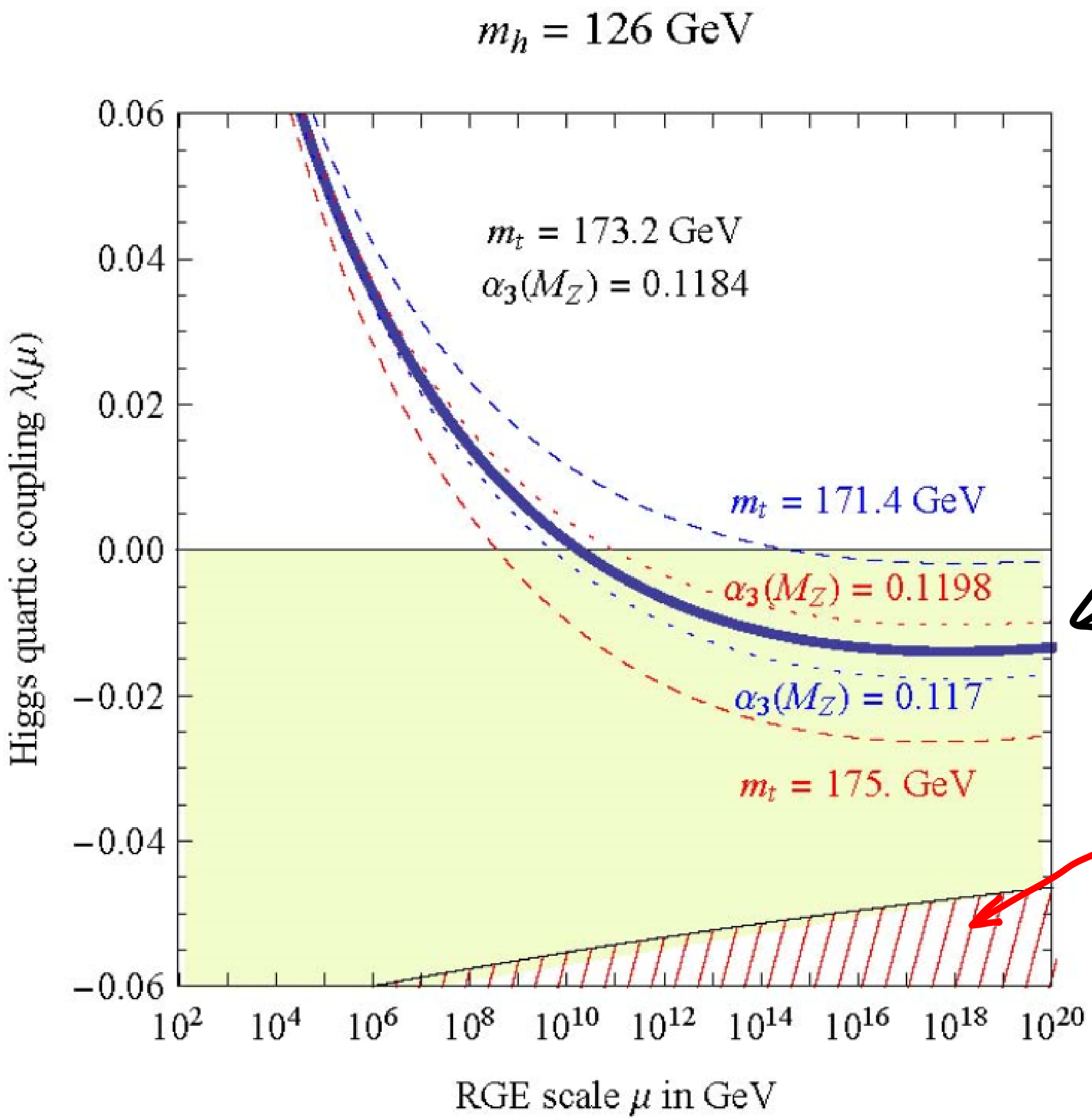
easily wins over τ_U^4

$p \lll 1$: Lifetime of EW vacuum much longer than τ_U

PROBABILITY OF VACUUM DECAY



LIFE IN A METASTABLE VACUUM

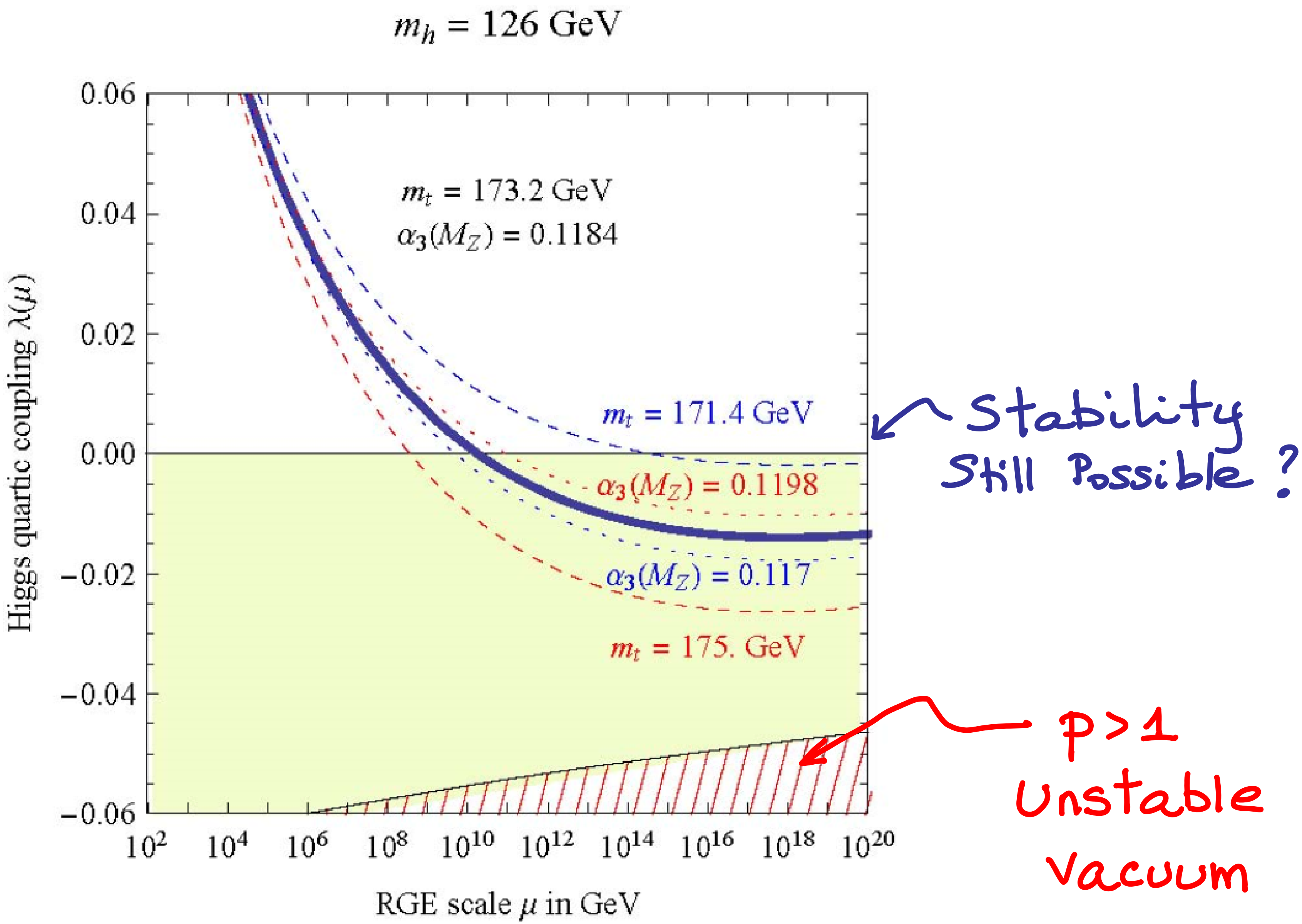


Lifetime $\propto \exp \frac{1}{|\lambda|}$

$\gg$ age of Universe

$p > 1$
Unstable vacuum

LIFE IN A METASTABLE VACUUM



NNLO STABILITY BOUND

Lower bound on M_h for stability up to M_{Pl}:

State-of-the-art NNLO calculation:

- 2-loop V_{eff} (Ford, Jack, Jones [ph/0111190])

- 3-loop RGES (..., Chetyrkin, Zoller [ph/1205.2892], Bednyakov, Pikelner, Velizhanin [ph/1212.6829])

- 2-loop matching in $\lambda \leftrightarrow M_h^2$; $h_t \leftrightarrow M_t$
  (..., Shaposhnikov et al [ph/1205.2893], , Degrassi et al [ph/1205.6497], , Buttazzo et al [ph/1307.3536])

NNLO STABILITY BOUND

For stability up to M_{Pl} :

$$M_h\,[\mathrm{GeV}] > 129.4 + 1.4\left(\frac{M_t(\mathrm{GeV}) - 173.1}{0.7}\right) - 0.5\left(\frac{\alpha_s(M_Z) - 0.1184}{0.0007}\right) \pm 1.0_{th}$$

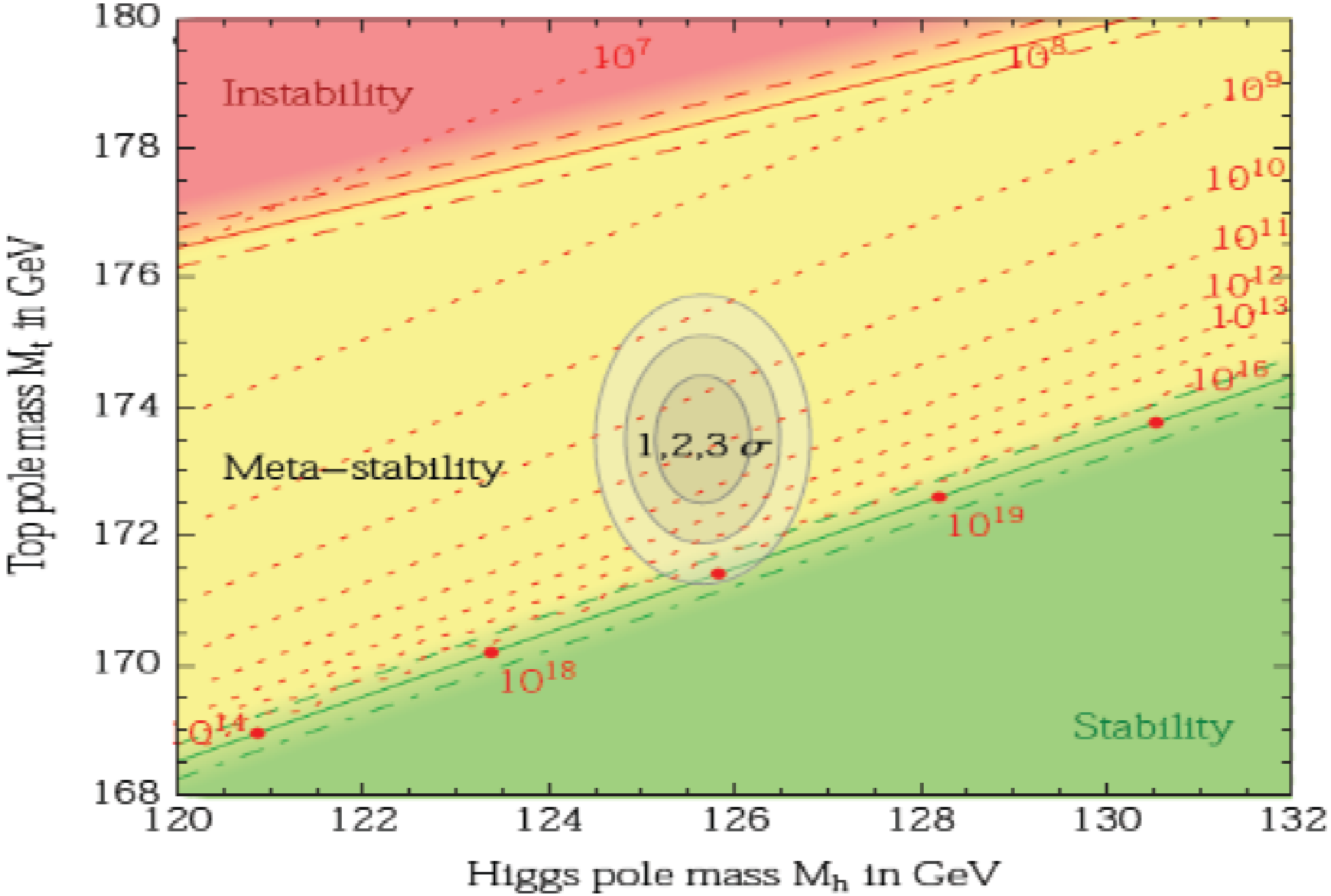
Degrassi et al '12

$$M_h\,[\mathrm{GeV}] > 129.6 + 2\left(\frac{M_t(\mathrm{GeV}) - 173.35}{1}\right) - 0.5\left(\frac{\alpha_s(M_Z) - 0.1184}{0.0007}\right) \pm 0.3_{th}$$

Buttazzo et al '13

Both reduced previous theory error by a factor 3

LIVING AT THE EDGE



TOP MASS CAVEATS

Have assumed

$$M_t = 173.1 \pm 0.7 \text{ GeV}$$

from Tevatron + LHC is the top pole mass.

Theoretically cleaner determination from $\sigma(t\bar{t})$ but larger error

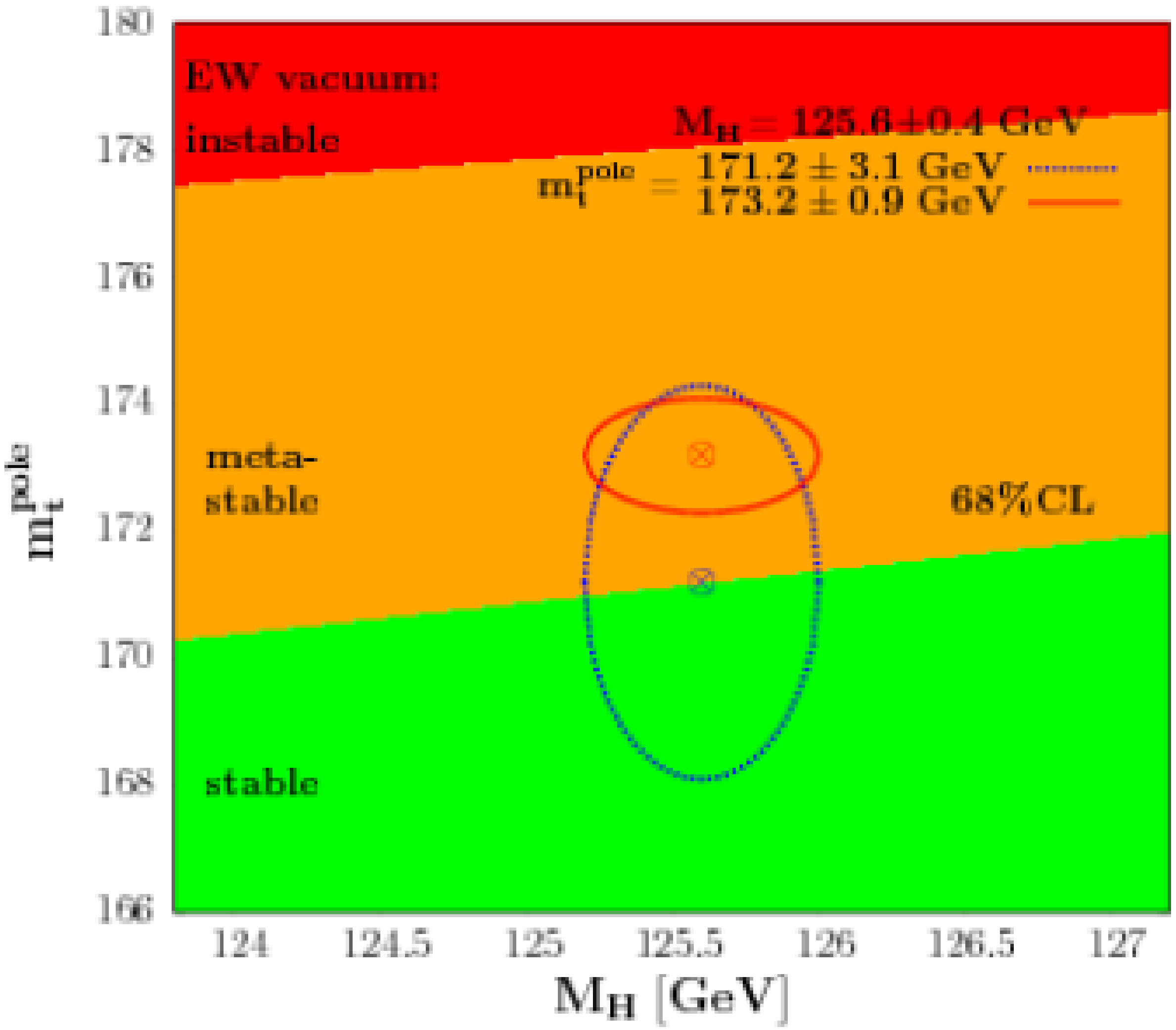
$$M_t = 171.2 \pm 3.1 \text{ GeV}$$

would still allow for stability

Alekhin, Djouadi, Moch '12

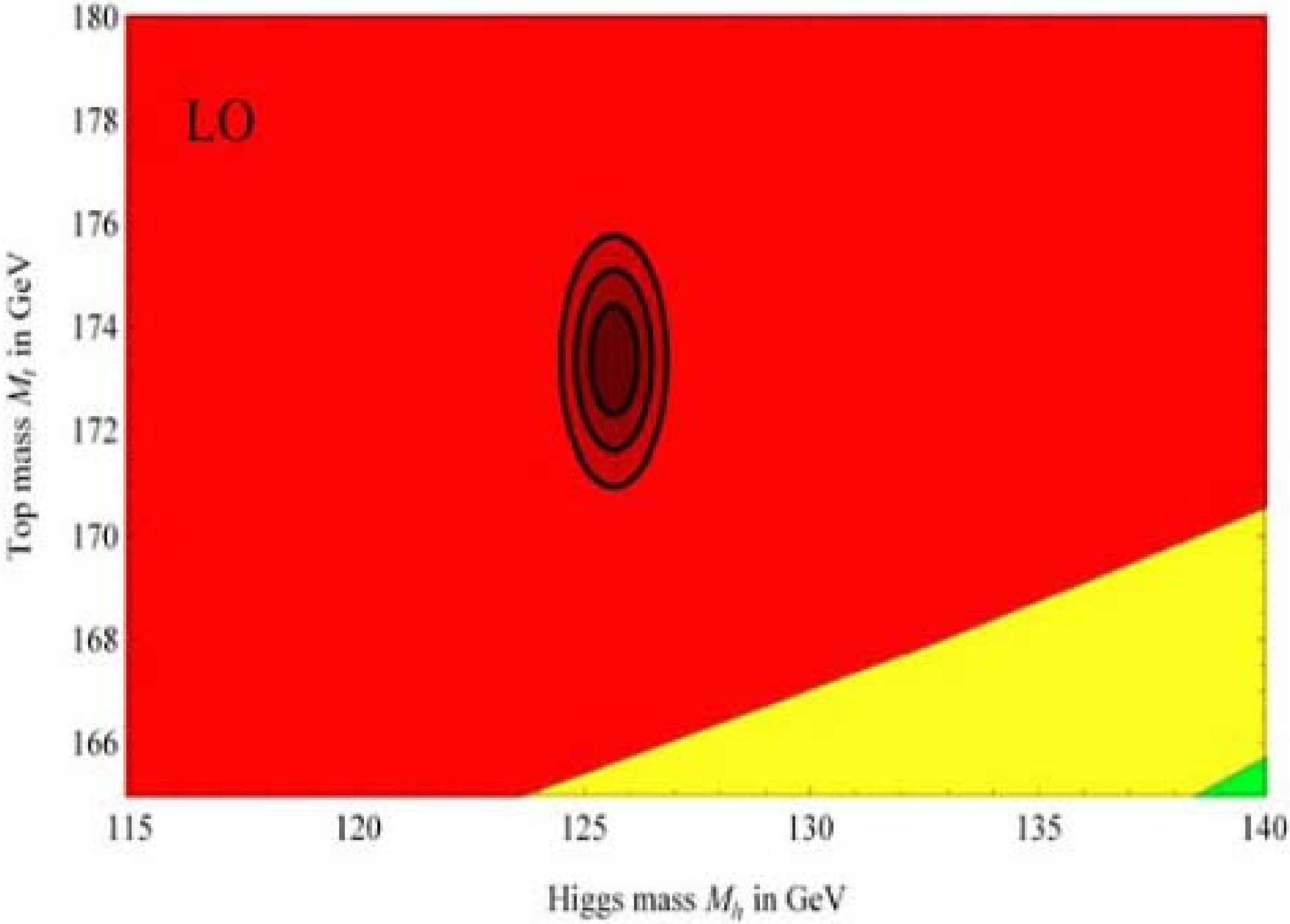
Too conservative given the good agreement...

TOP MASS CAVEATS



Alekhin, Djouadi, Moch'13

PROGRESS IN STABILITY CALCULATION



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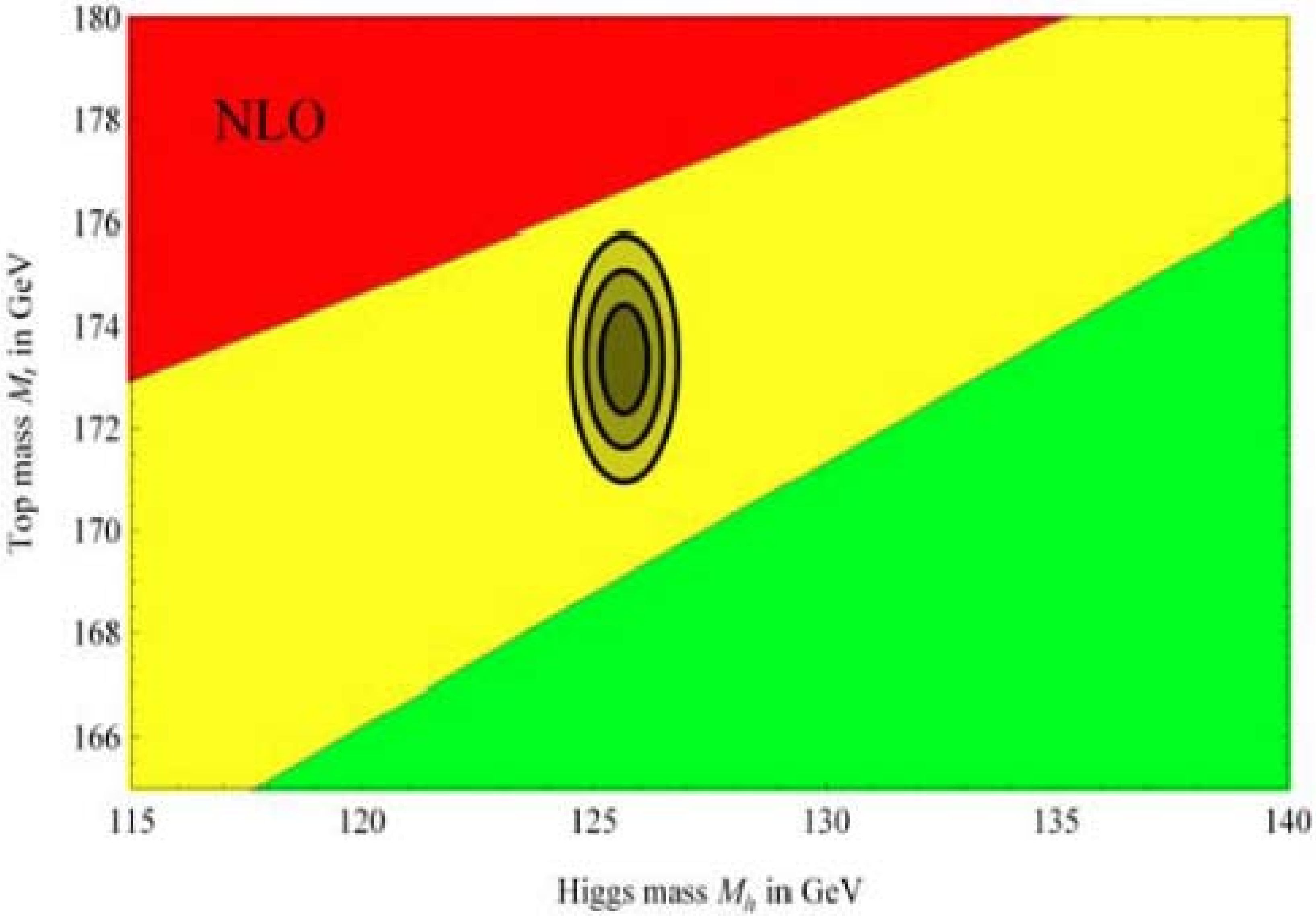

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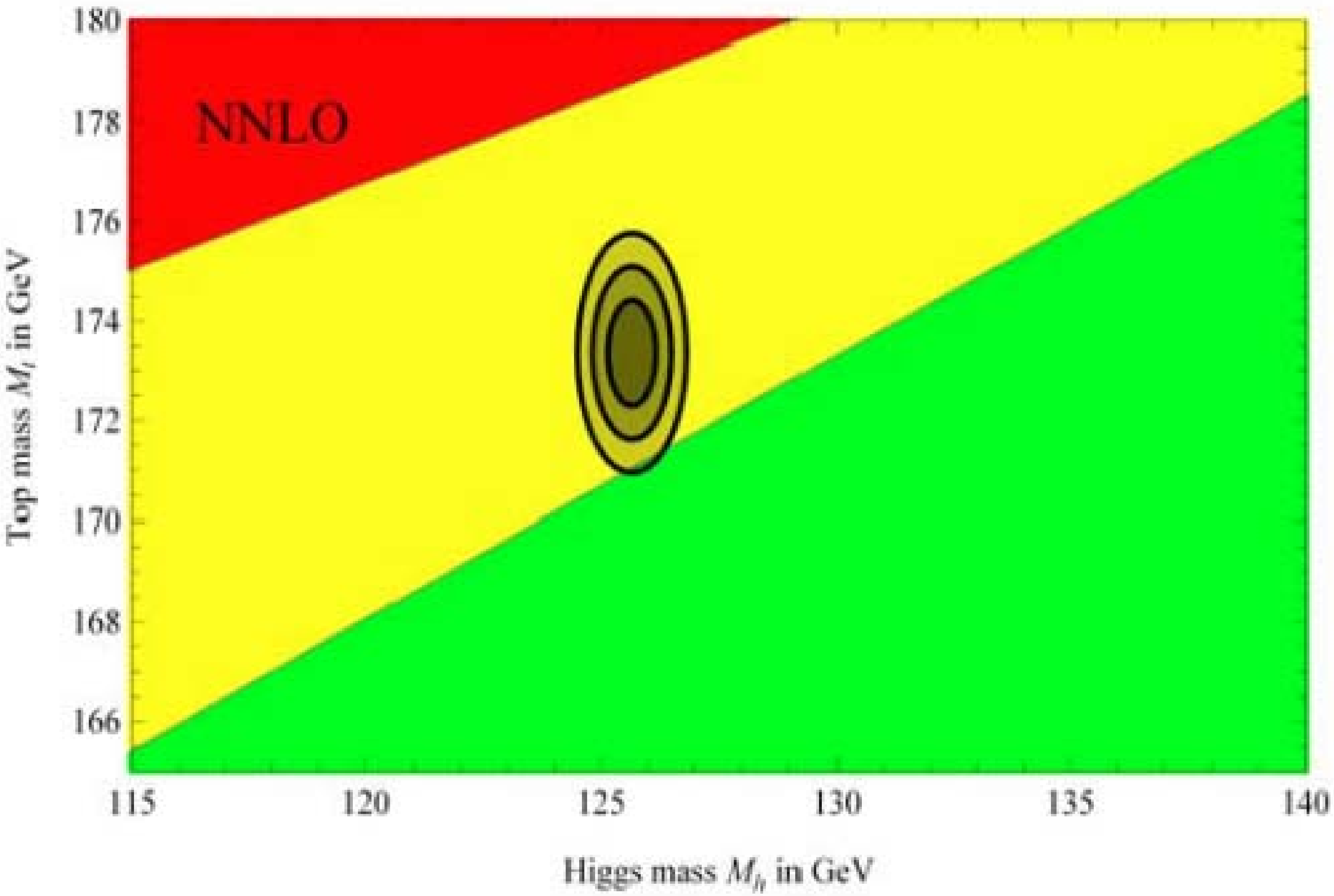
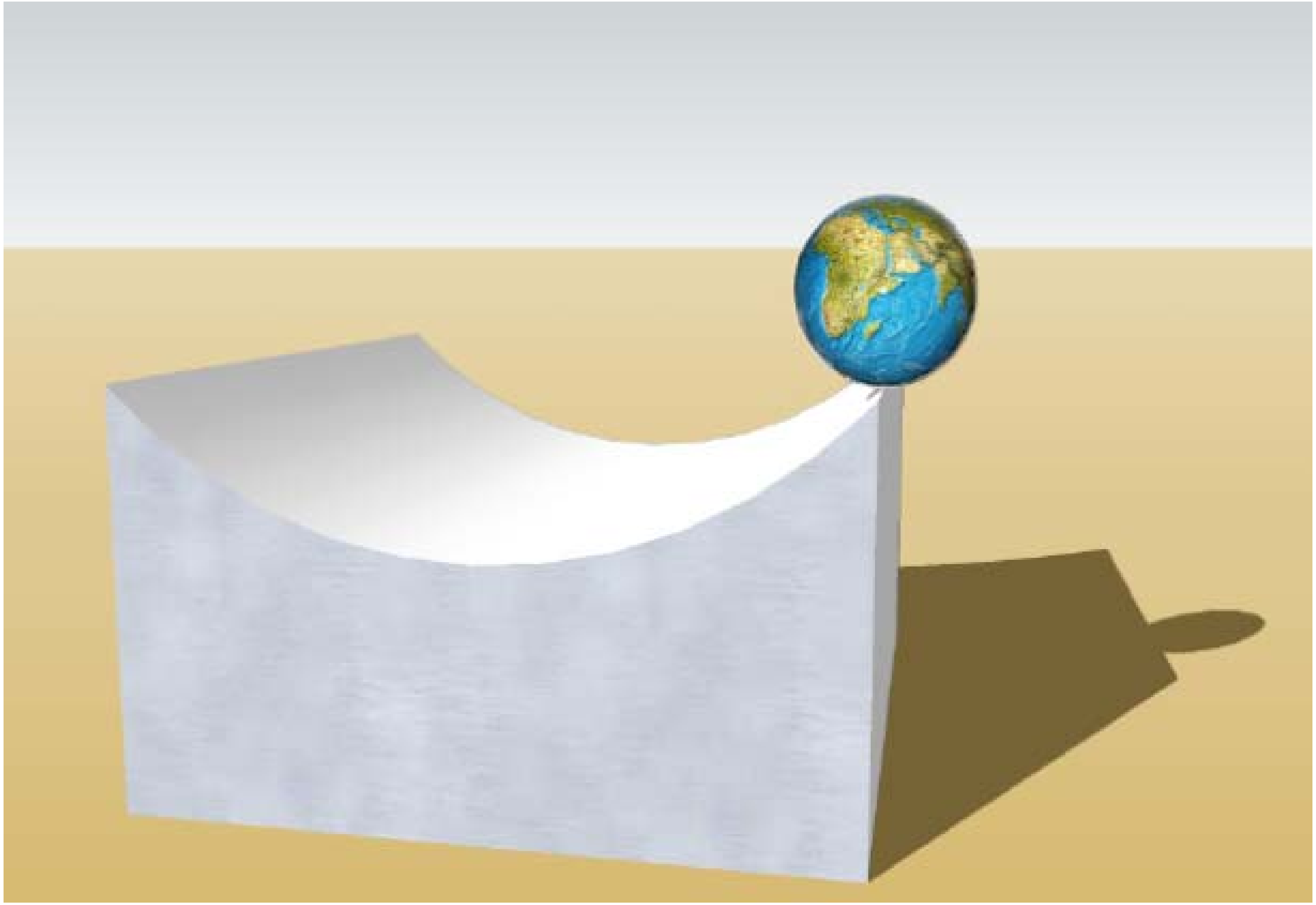


LIVING AT THE EDGE



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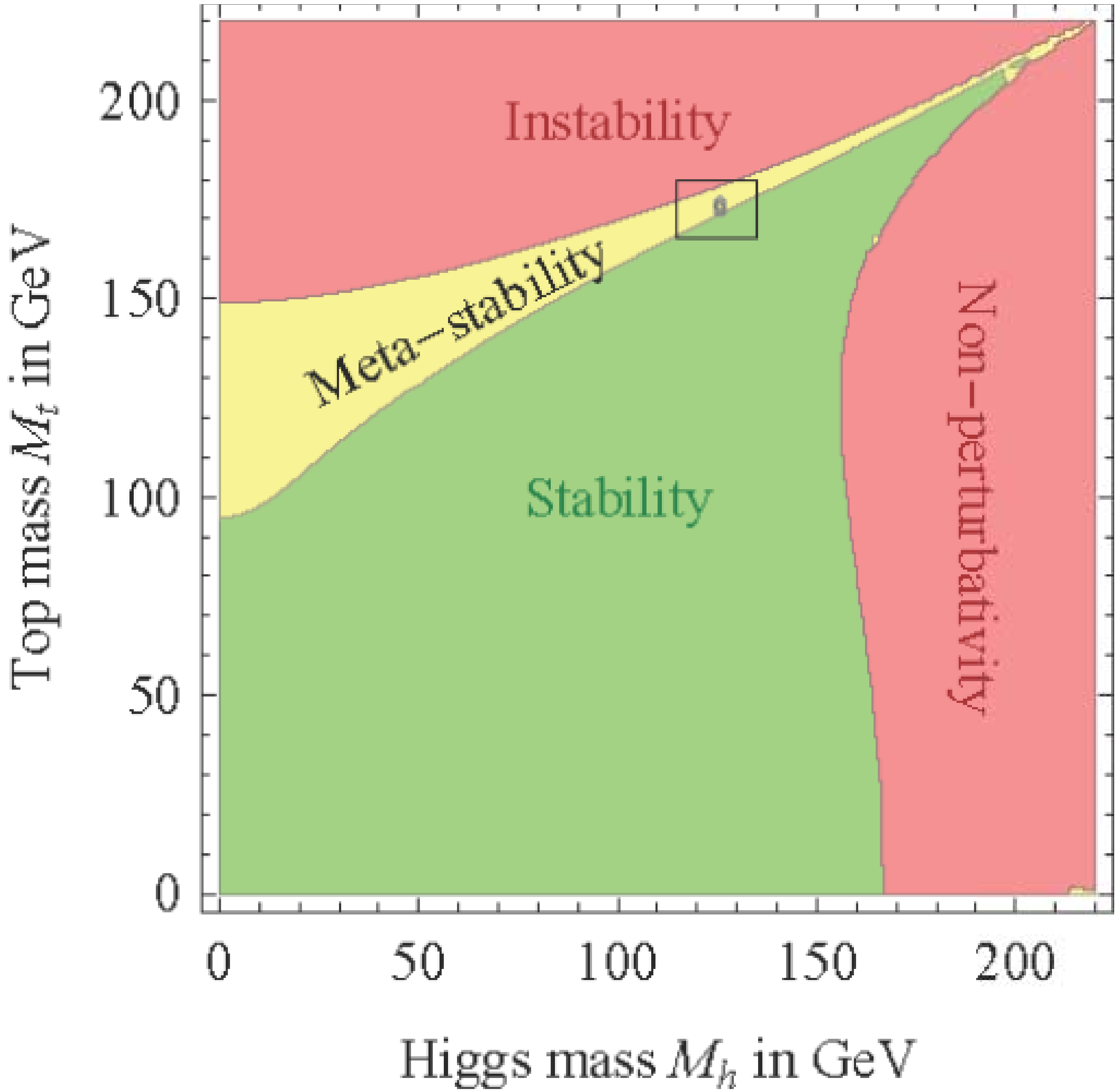
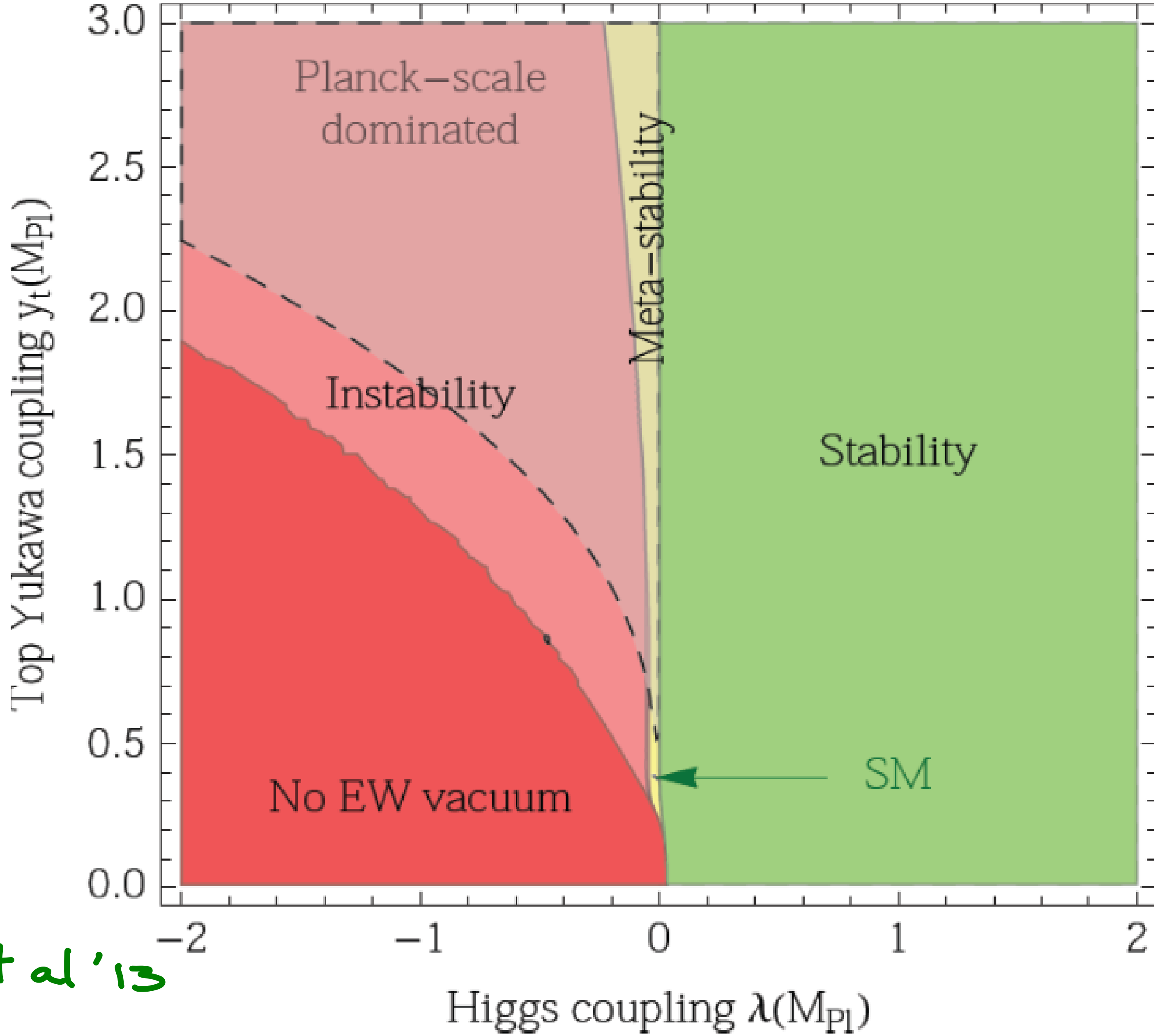


LIVING AT THE EDGE



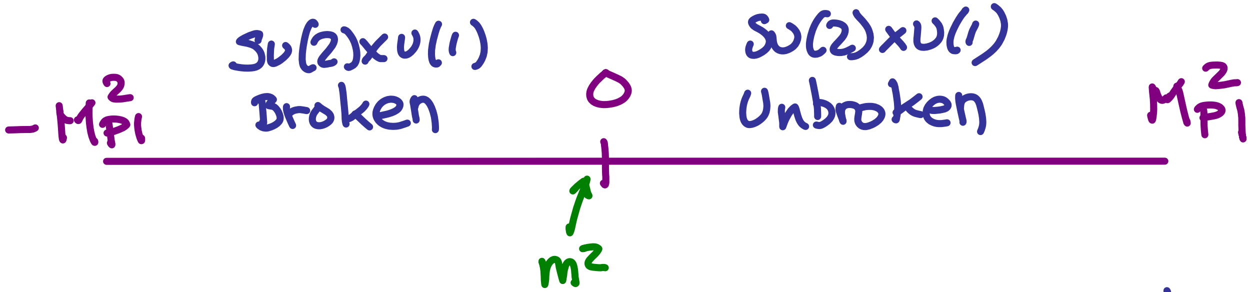
Buttazzo et al '13

NEW KNOWLEDGE BRINGS NEW QUESTIONS

- Why do we live near the critical boundary for stability?

$$\lambda(M_{Pl}) \simeq 0$$

- Is this related to our living near the phase boundary $m^2/M_{Pl}^2 \simeq 0$?



- Is the EW scale determined by Planck scale physics?
- Or is this just a coincidence? BSM...

BSM & STABILITY

Even without naturalness, BSM must exist...

Its impact on the Higgs instability can be

IRRELEVANT

MAKE IT WORSE

CURE IT

BSM & STABILITY

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Its impact on the Higgs instability can be

Example

IRRELEVANT See-saw neutrinos

MAKE IT WORSE

CURE IT

BSM & STABILITY

Even without naturalness, BSM must exist...

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	Example
IRRELEVANT	See-saw neutrinos
MAKE IT WORSE	See-saw neutrinos
CURE IT	

BSM & STABILITY

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IRRELEVANT	See-saw neutrinos
MAKE IT WORSE	See-saw neutrinos
CURE IT	See-saw neutrinos

BSM & STABILITY

Even without naturalness, BSM must exist...

Its impact on the Higgs instability can be

	Example
IRRELEVANT	See-saw neutrinos $M_R \lesssim 10^{13}$ GeV
MAKE IT WORSE	See-saw neutrinos $M_R \gtrsim 10^{13}$ GeV
CURE IT	See-saw neutrinos $M_R \sim \langle S \rangle$ & $\lambda_{HS} \lvert H \rvert^2 \lvert S \rvert^2$

Lebedev '12, Elias-Miro et al. '12

OTHER IMPLICATIONS

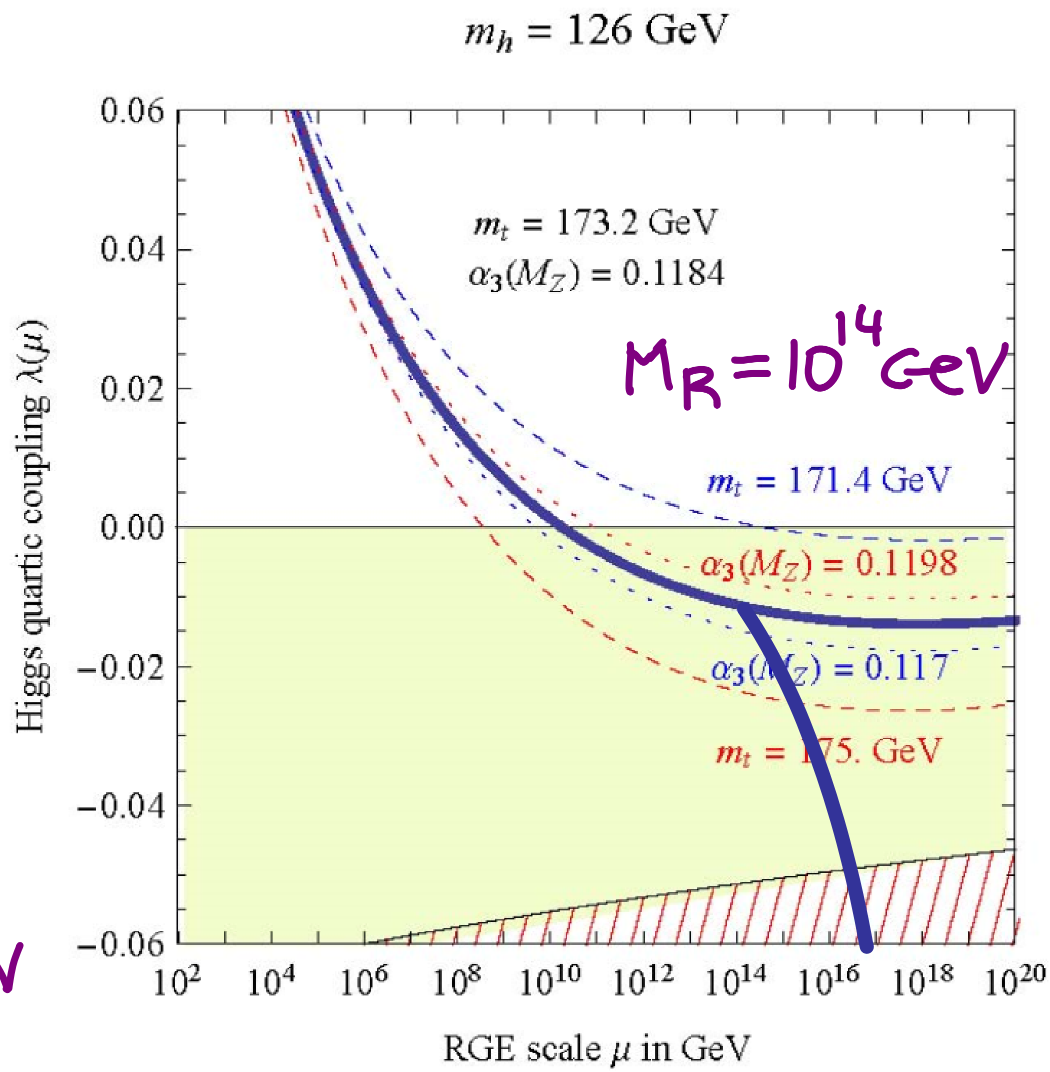
- See-saw neutrinos : Impact on $\beta_\lambda \supset -y_\nu^4/(16\pi^2)$ *

$$m_\nu \sim \frac{y_\nu^2 v^2}{M_R}$$

$M_R \uparrow \Rightarrow y_\nu \uparrow$

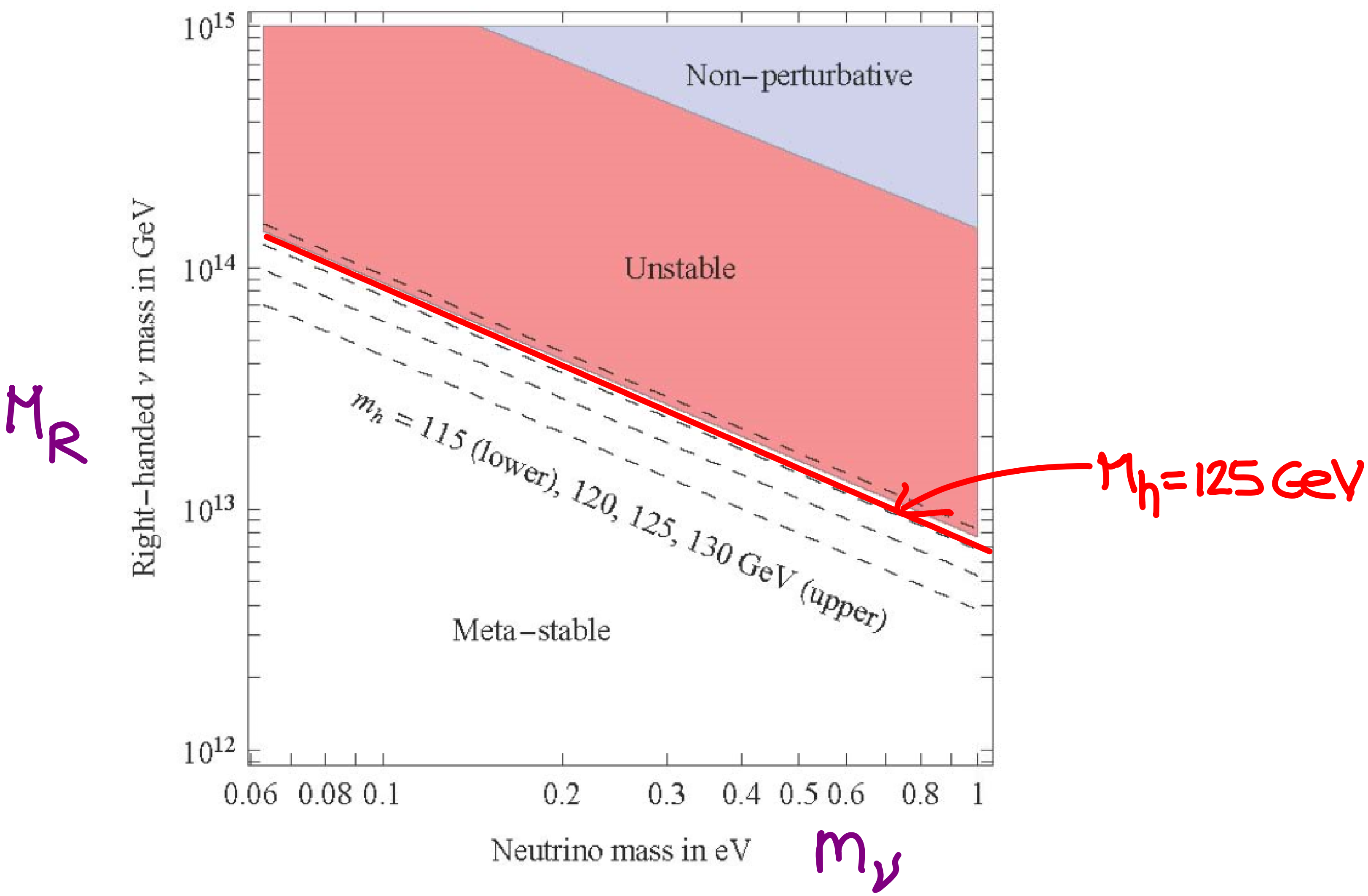
Adds to the top destabilizing effect

Important for $M_R \gtrsim 10^{13-14}$ GeV



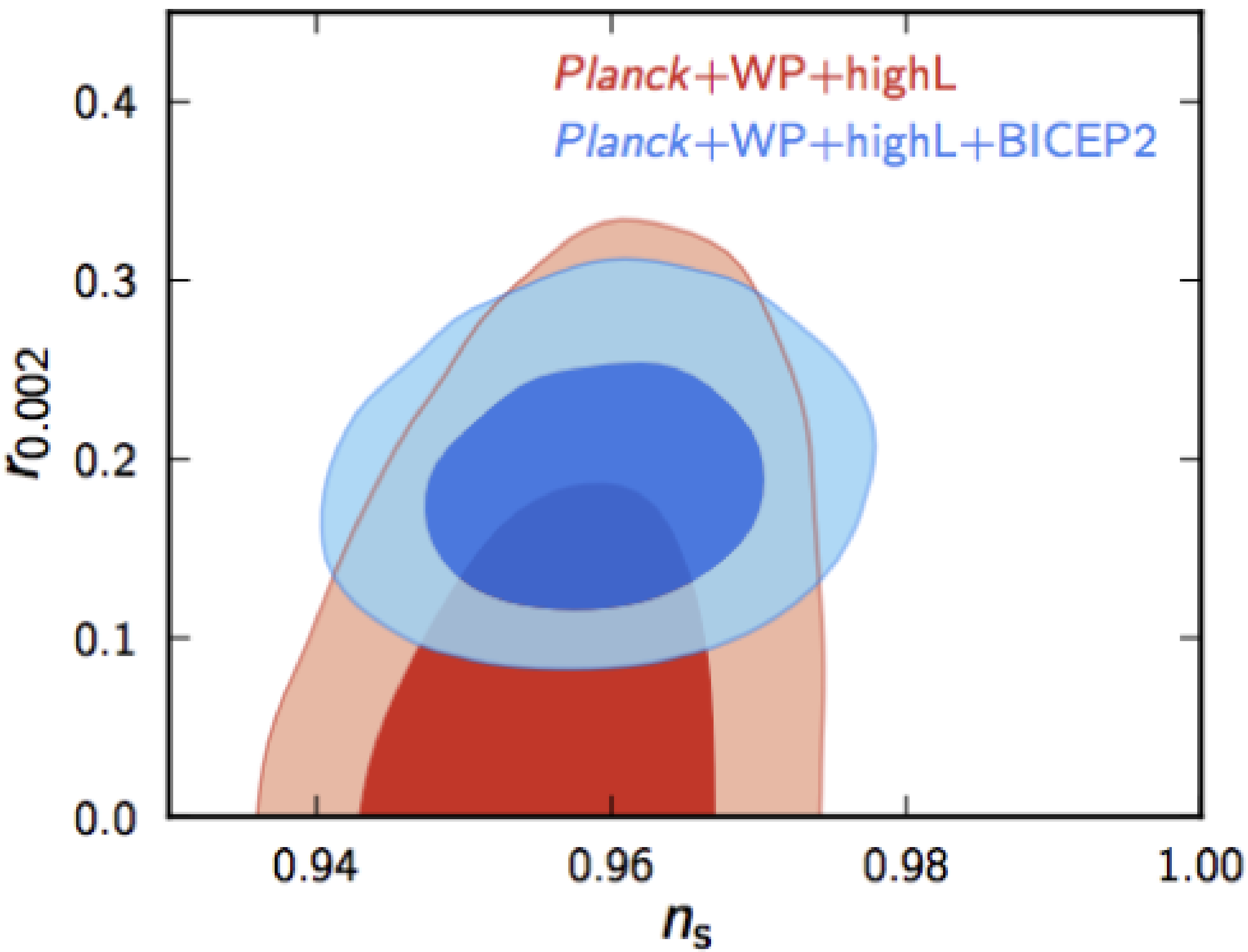
OTHER IMPLICATIONS

- See-saw neutrinos: Bound on M_{ν_R}



INTERPLAY WITH COSMOLOGY

BICEP2 big news from Monday:



$\Rightarrow U_I \sim (10^{16}\,\mathrm{GeV})^4$

$= \dfrac{H_I^2}{3M_P^2}$

$\Downarrow$

$H_I \simeq 10^{13}\,\mathrm{GeV}$

Dangerous for stability

INTERPLAY WITH COSMOLOGY

Inflation causes light fields to fluctuate

Long-wavelength modes $\sim$ homogeneous classical field

$\langle h \rangle \sim H_I > \Lambda_I^{1/2} \Rightarrow$ Vacuum decay

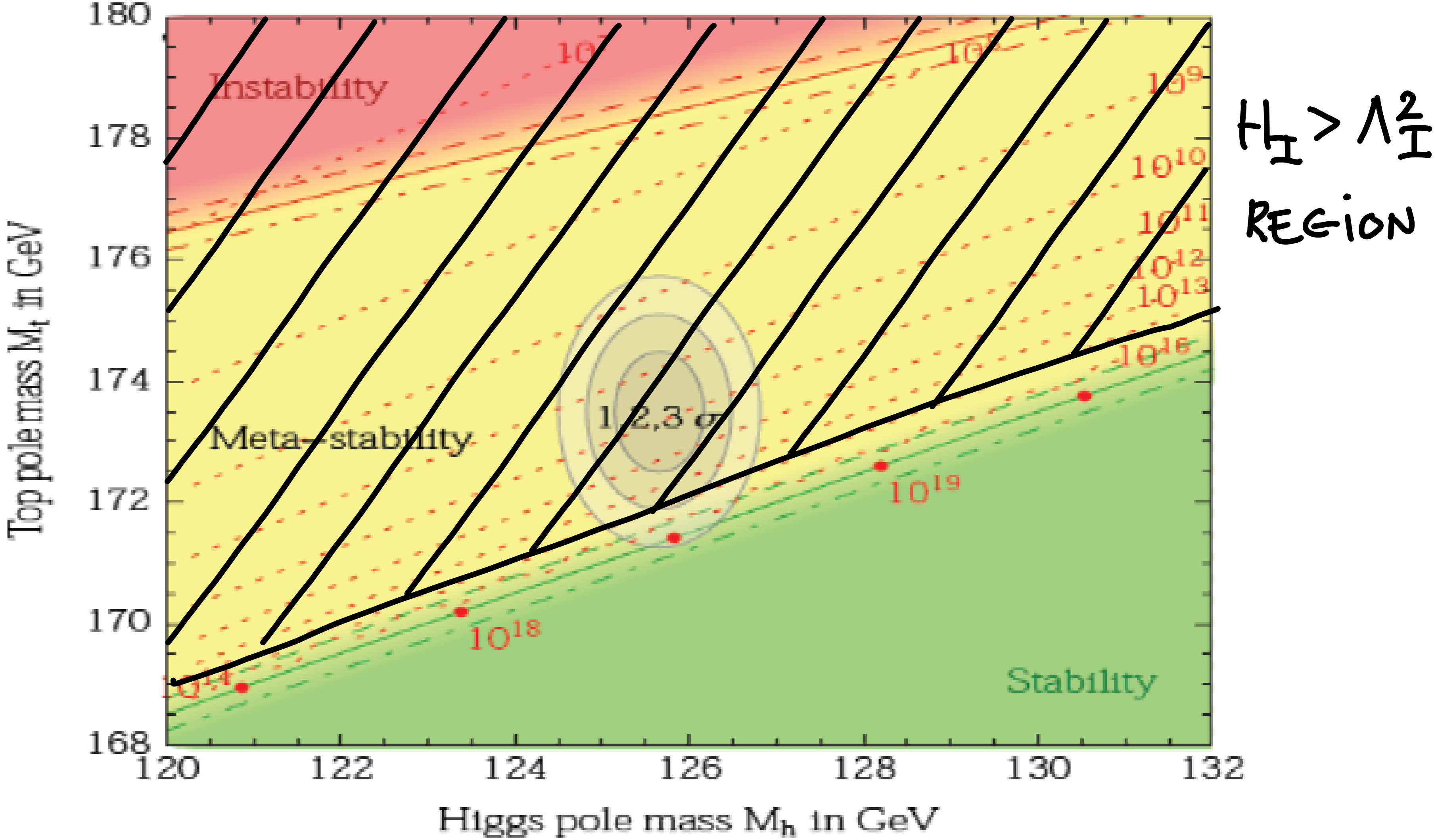
Survival probability

$$\mathcal{P}_s \sim \exp\left(- \frac{H_I N_e}{32 \Lambda_I^2} \right)$$

Exponentially suppressed for $H_I \gg \Lambda_I$

Our universe would be very unlikely.

INTERPLAY WITH COSMOLOGY



CONCLUSIONS

We finally have data to explore the physics of electroweak symmetry breaking!

★ $M_h \simeq 125\,\mathrm{GeV}$

⇨ Unstable EW vacuum in SM ($\Lambda_I \sim 10^{10}\,\mathrm{GeV}$)

EW vacuum is Long-lived and intriguinly close to stability boundary

Deep meaning of this ?

This instability has implications for BSM, cosmology...

Further motivations to improve the determination of M_{top}

NNLO INGREDIENTS

Renormalisation Group Equations

	LO 1 loop	NLO 2 loop	NNLO 3 loop	NNNLO 4 loop
g_3	full [53, 54]	$\mathcal{O}(\alpha_3^2)$ [55, 56] $\mathcal{O}(\alpha_3\alpha_{1,2})$ [61] full [63]	$\mathcal{O}(\alpha_3^3)$ [57, 58] $\mathcal{O}(\alpha_3^2\alpha_t)$ [62] full [64, 65]	$\mathcal{O}(\alpha_3^4)$ [59, 60]
$g_{1,2}$	full [53, 54]	full [63]	full [64, 65]	—
y_t	full [66]	$\mathcal{O}(\alpha_t^2, \alpha_3\alpha_t)$ [67] full [70]	full [68, 69]	—
λ, m^2	full [66]	full [71, 72]	full [73, 74]	—

Threshold corrections at the weak scale

	LO 0 loop	NLO 1 loop	NNLO 2 loop	NNNLO 3 loop
g_2	$2M_W/V$	full [75, 76]	full [This work]	—
g_Y	$2\sqrt{M_Z^2 - M_W^2}/V$	full [75, 76]	full [This work]	—
y_t	$\sqrt{2}M_t/V$	$\mathcal{O}(\alpha_3)$ [77] $\mathcal{O}(\alpha)$ [81]	$\mathcal{O}(\alpha_3^2, \alpha_3\alpha_{1,2})$ [34] full [This work]	$\mathcal{O}(\alpha_3^3)$ [78–80]
λ	$M_h^2/2V^2$	full [82]	for $g_{1,2} = 0$ [4] full [This work]	—
m^2	M_h^2	full [82]	full [This work]	—

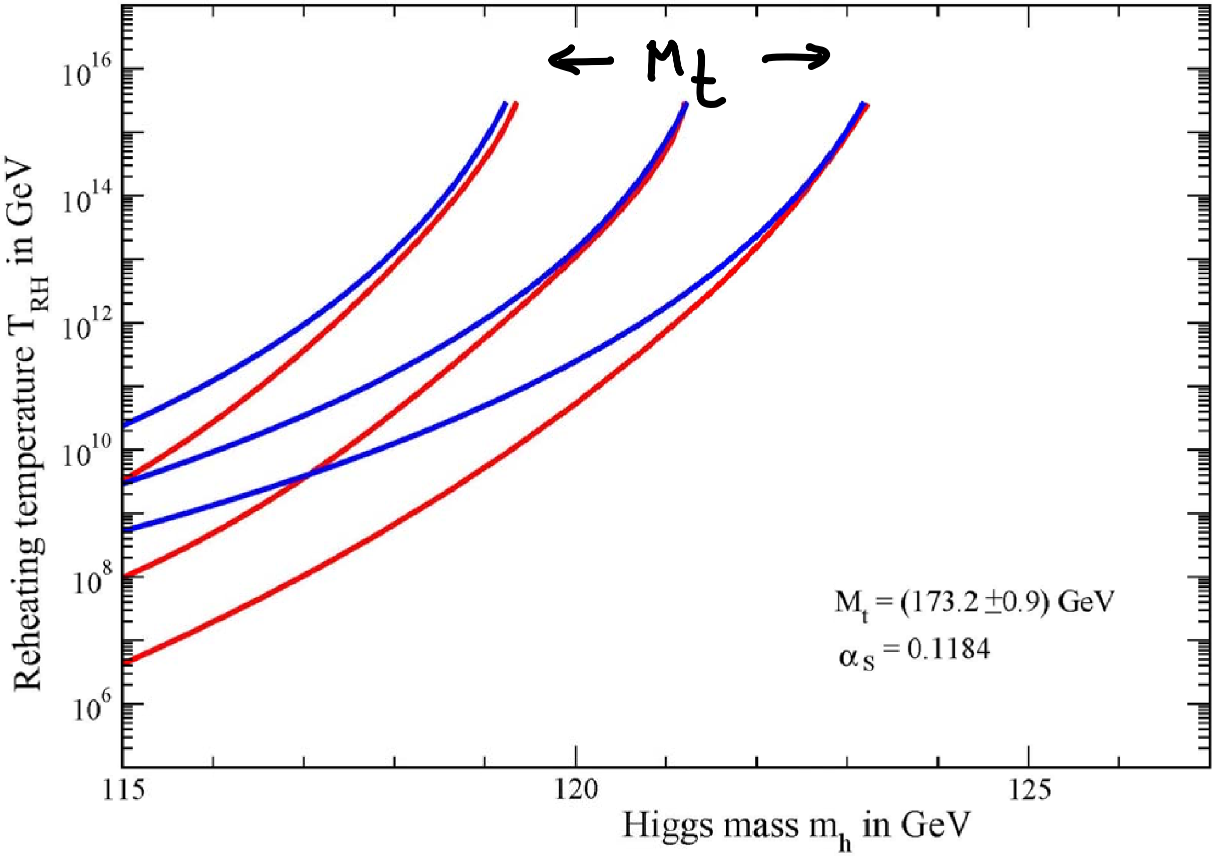
Table 1: *Present status of higher-order computations included in our code. With the present paper the calculation of the SM parameters at NNLO precision is complete. Here we have defined* $V \equiv (\sqrt{2}G_\mu)^{-1/2}$ *and* $g_1 = \sqrt{5/3}g_Y$.

Buttazzo et al '13.

OTHER IMPLICATIONS

- Cosmology :

  Thermal fluctuations can induce vacuum decay



Bound on T_{RH} ?

OTHER IMPLICATIONS

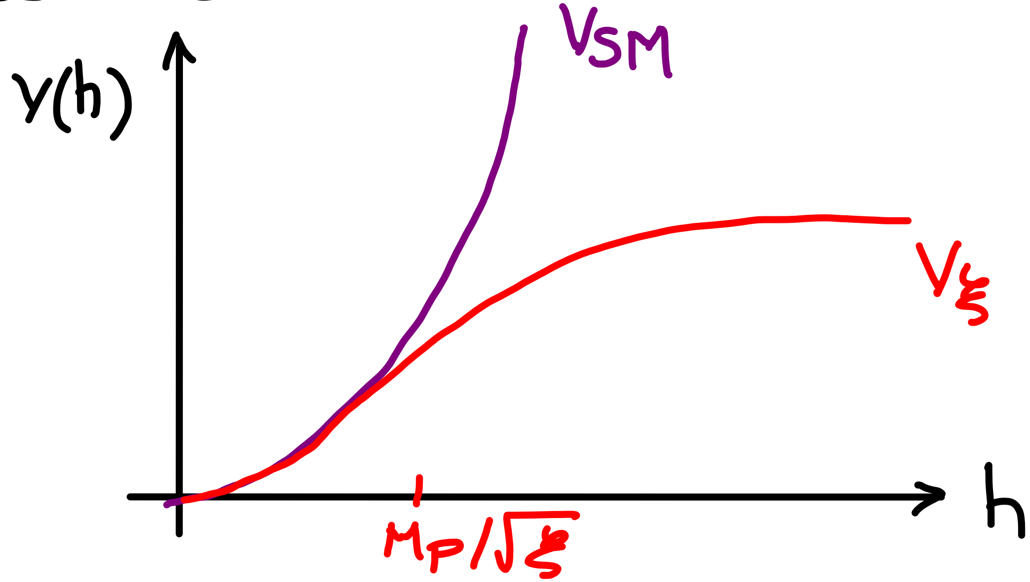
- Cosmology: Higgs inflation Bezrukov, Shaposhnikov '07

Higgs coupled to gravity as $\mathcal{L} \supset \int \sqrt{-g}\, \xi\, |H|^2 R$

coupling removed by $g_{\mu\nu} \to g_{\mu\nu} (1 + \xi h^2/M_P^2)^{-1}$

rescales the potential as

$$V(h) \Rightarrow \frac{V(h)}{(1 + \xi h^2/M_P^2)^2}$$



Requires $\xi \sim 10^4$ to give the right spectrum of primordial fluctuations.

(MORE) TROUBLE FOR HIGGS INFLATION

※1 Effective theory with cutoff

$$\Lambda \sim \frac{M_P}{\xi} \ll \Lambda_{HI} \sim \frac{M_P}{\sqrt{\xi}}$$

Can't trust the plateau region

Burgess, Lee, Trott '09. Barbón, JRE '09

※2 Stability up to $\sim 10\Lambda_{HI}$ is a must.

Requires marginal values of M_h & M_t