Mechanical quantum sensing in the search for new physics

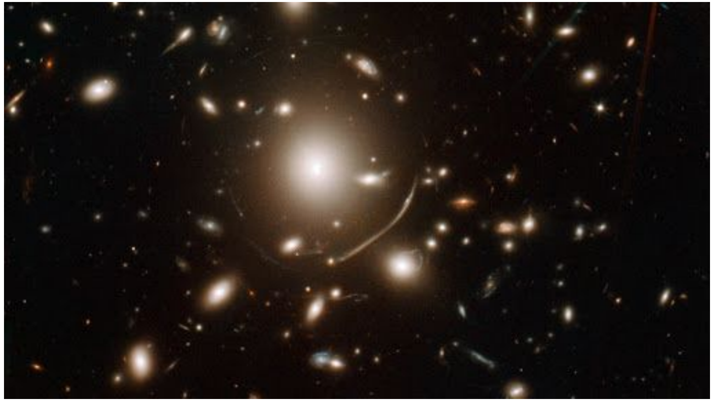
Daniel Carney

BERKELEY LAB

Assuming dark matter exists in the first place (!), the *only coupling it is guaranteed to have to visible matter* is through gravity.

Local dark matter density ~ one proton mass per cm^3. Hopeless to try to detect it through this gravitational force in a local lab. Right?

Extremely hard, but maybe possible...



1605.01103 Hall, Callister, Frolov, Muller, Pospelov, Adhikari
1809.00968 Kawasaki
1903.00492 Carney, Ghosh, Krnjaic, Taylor

How to view this talk

Ok, so if we can just detect the vacuum fluctuations of some 40kg mirrors separated by a few km, we can detect gravitational waves… assuming the sources exist in the right mass range. What? That’s seven orders of magnitude better sensitivity than what anybody can do now? Relax, we’ll figure it out...

Contents

- Detection problem
- Where we are now, and what's between here and gravitational detection
- A brief meditation on the word "possible"

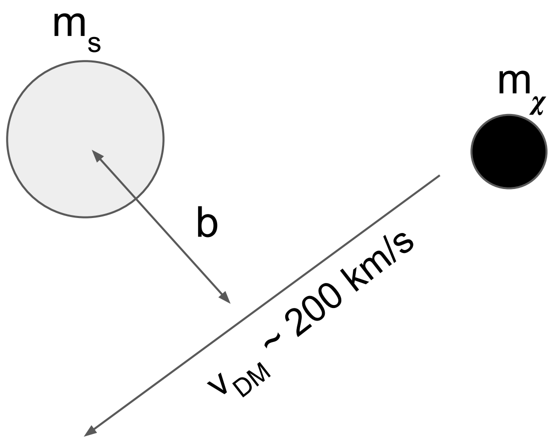
The basic idea

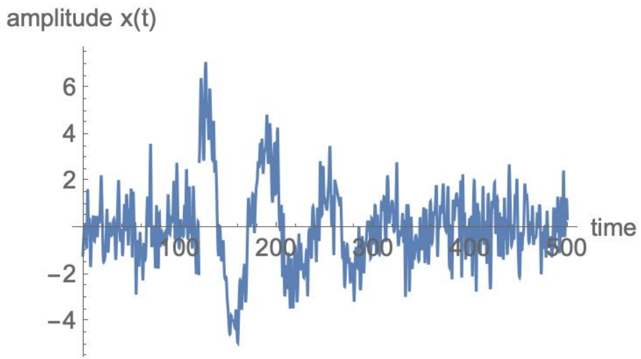
$$F = \frac{G_N m_s m_\chi}{r^2}$$

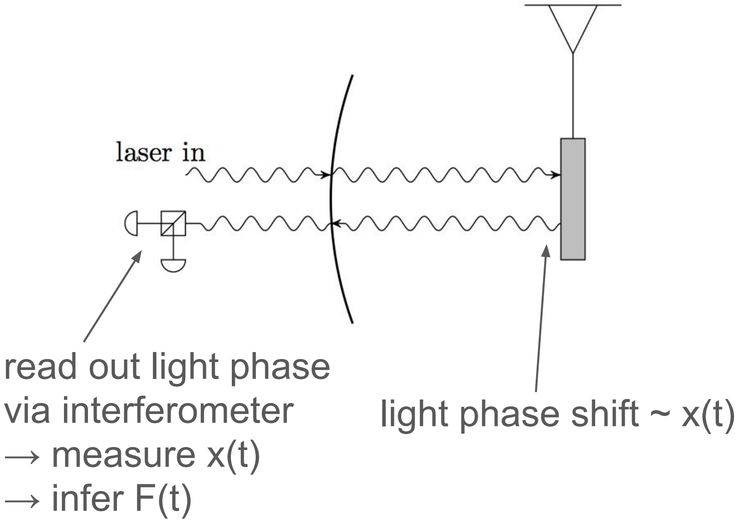
→ want heavy DM, small impact parameter, very good force sensor

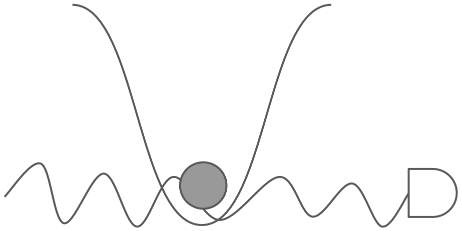
$$R = \frac{\rho_\chi v A_d}{m_\chi} \sim \frac{1}{\text{year}} \left(\frac{m_{\text{Pl}}}{m_\chi}\right) \left(\frac{A_d}{1\,\text{m}^2}\right)$$

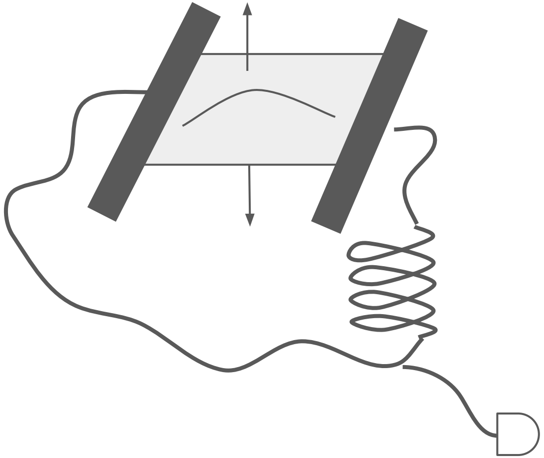
→ want large area

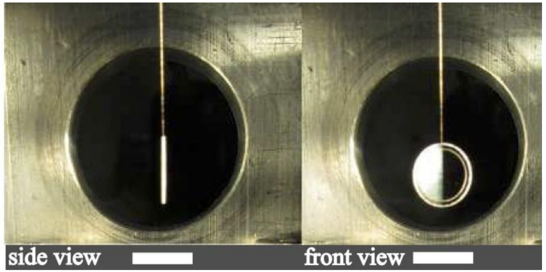




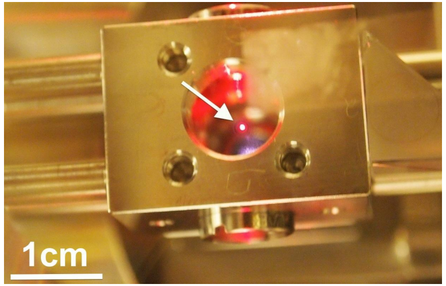




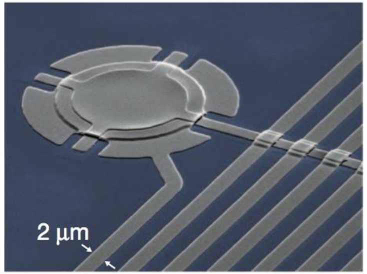




Matsumoto et al, PRA 2015



Moore group @ Yale



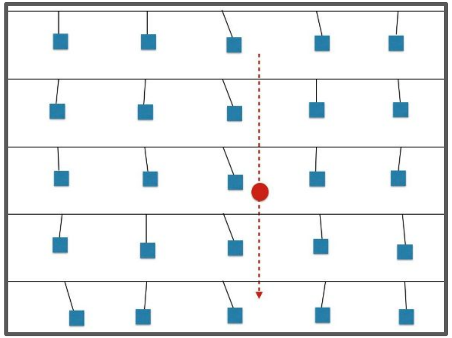
Teufel et al, Nature 2011

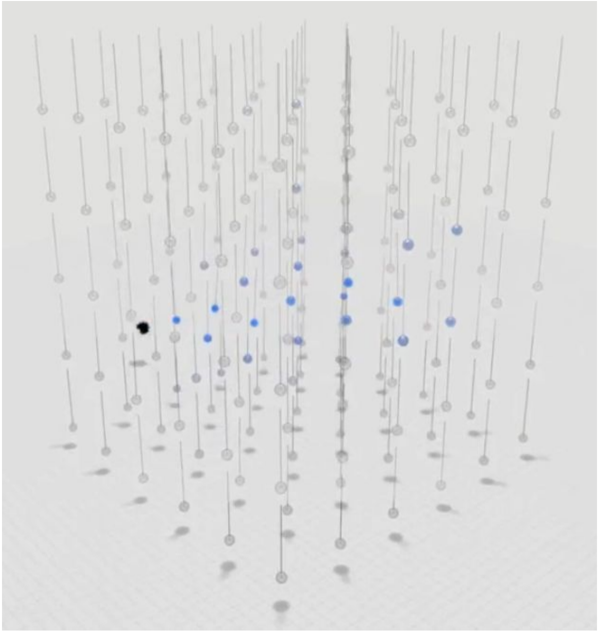
Windchime array concept

Signal = correlated track of sensors moving

- Complete directional info
- Exquisite background rejection
- sqrt(N) noise reduction (N = sensors on track)

Scales to keep in mind: ~ mm-cm spacing, mg-g mass devices, the more detectors the better (more on these numbers later)





PHYSICAL REVIEW LETTERS

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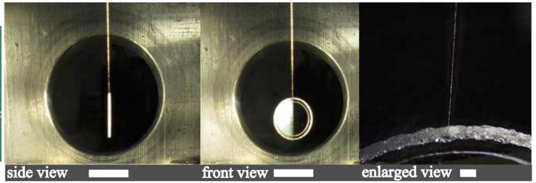

Featured in Physics

Demonstration of Displacement Sensing of a mg-Scale Pendulum for mm- and mg-Scale Gravity Measurements

Nobuyuki Matsumoto, Seth B. Cataño-Lopez, Masakazu Sugawara, Seiya Suzuki, Naofumi Abe, Kentaro Komori, Yuta Michimura, Yoichi Aso, and Keiichi Edamatsu

Phys. Rev. Lett. **122**, 071101 – Published 19 February 2019

Physics See Synopsis: Gravity of the Ultralight

nature

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Article | Published: 10 March 2021

Measurement of gravitational coupling between millimetre-sized masses

Tobias Westphal, Hans Hepach, Jeremias Pfaff & Markus Aspelmeyer

Nature **591**, 225–228 (2021) | Cite this article

19k Accesses | **18** Citations | **673** Altmetric | Metrics

$F_{grav} = G_N\, m^2/d^2 \sim 10^{-17}$ N

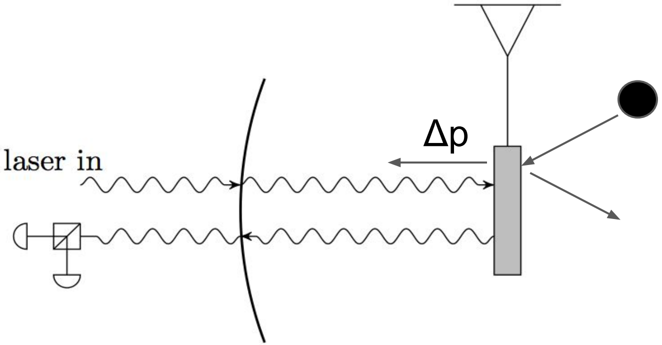
Note the conversion factor $m_{planck} = 0.02$ mg

Our problem is harder: 200 km/s DM velocity → only have ~ns-us time to integrate the signal

Noise

Photo: Hiroshi Sugimoto

Impulse sensing problem

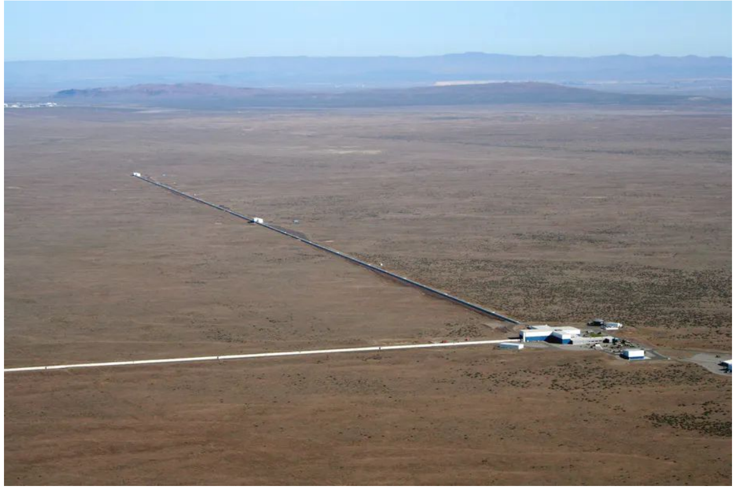


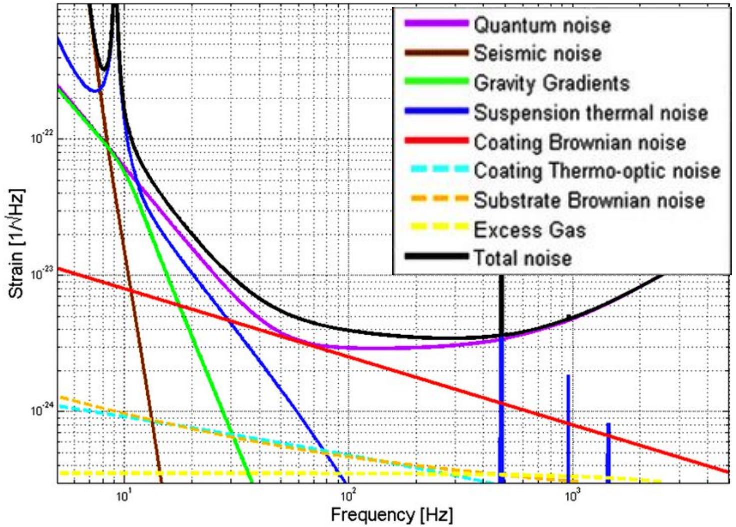
Goal: detect a sharp impulse (= ∫Fdt = Δp).

With sufficient technical noise control (stray field reduction, etc.), any detector is ultimately limited by

- Thermal noise
- Quantum measurement noise

Quantum-limited detection





The Sensitivity of the Advanced LIGO Detectors at the Beginning of Gravitational Wave Astronomy
LIGO Collaboration 1604.00439

SNR at thermal noise level

If thermal noise dominates:

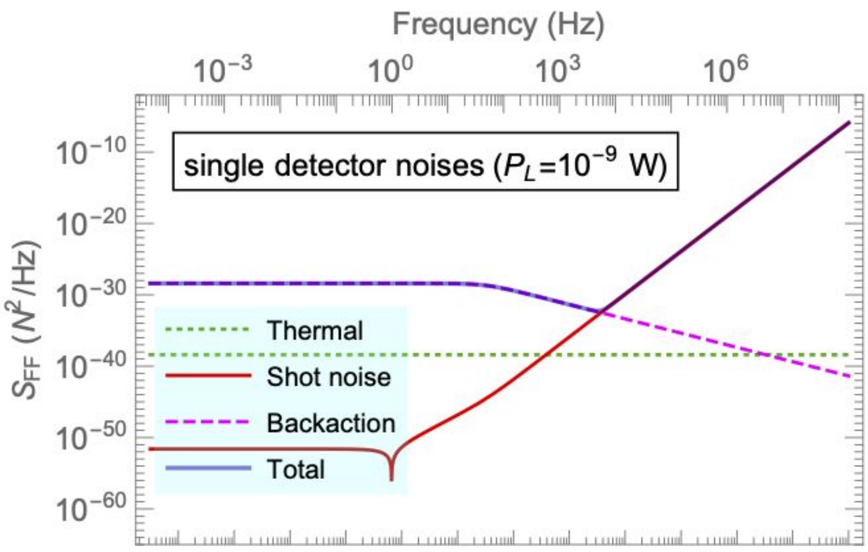
$$\mathrm{SNR}_{\mathrm{thermal}} = \frac{G_N m_\chi m_s / bv}{\sqrt{(4 m_s k_B T \omega / Q)(b/v)}}$$

$$\approx 0.5 \times \left(\frac{m_\chi}{1\ \mathrm{mg}}\right) \left(\frac{m_s}{1\ \mathrm{mg}}\right)^{1/2} \left(\frac{1\ \mathrm{mm}}{b}\right)^{3/2}$$

→ **Gravitational detection is possible if we can get to thermal noise floor**

But notice in the plot: measurement noise >> thermal.

→ **Need to reduce the quantum noise**



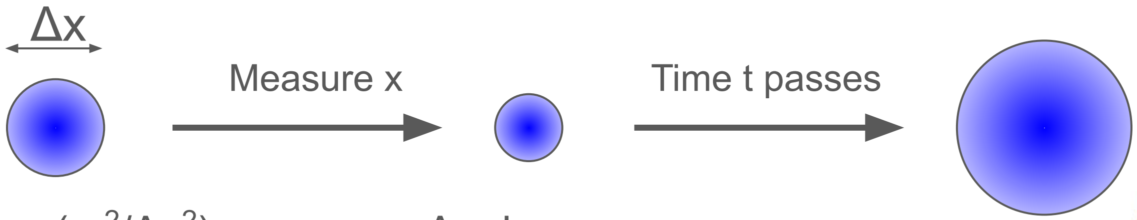
Sensor m = 1 mg, frequency = 1 Hz, in dil fridge

Quantum measurement noise

PHYSICAL REVIEW D VOLUME 23, NUMBER 8 15 APRIL 1981

Quantum-mechanical noise in an interferometer

Carlton M. Caves
W. K. Kellogg Radiation Laboratory, California Institute of Technology, Pasadena, California 91125
(Received 15 August 1980)



$\psi = \exp(-x^2/\Delta x^2)$

$\Delta x \Delta p = \hbar/2$
minimal uncertainty

Δx decreases
↓
Δp increases

$$\Delta x \rightarrow \Delta x + \frac{\Delta p}{2m}t$$

optimize → “standard quantum limit”

$$\Delta x_{SQL} = \sqrt{\frac{\hbar t}{m}}$$

Quantum limits in impulse sensing

Similarly there's an SQL for impulses:

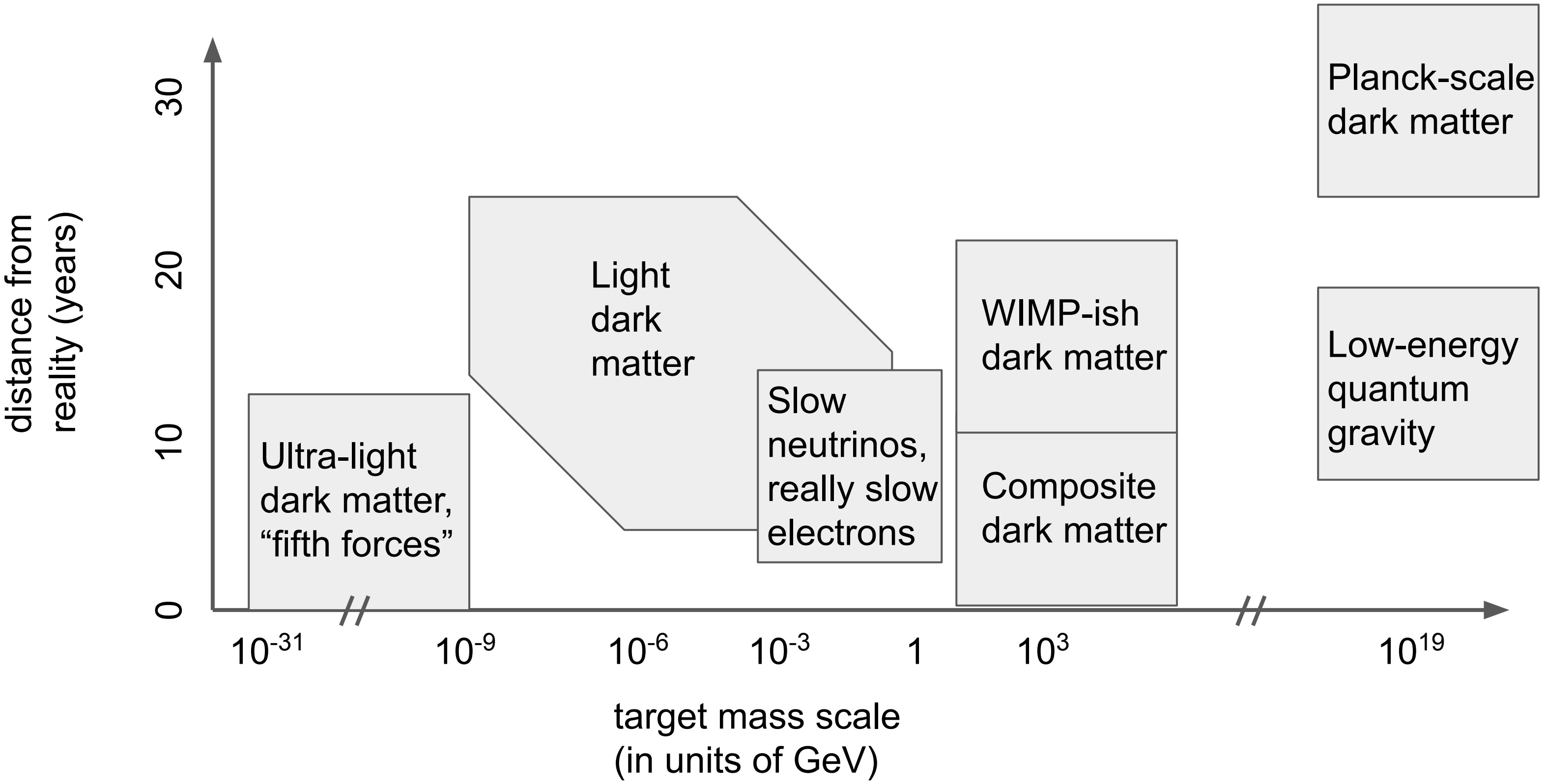
$$\Delta p_{SQL} = \sqrt{\hbar m_s \omega}$$

~ 600 keV (m = 1 ng, ω = 1 kHz)

Measurement at SQL means that you can resolve the motion of the sensor with error = sensor ground state wavefunction uncertainty.

Incredibly, we need to do at least 3-4 orders of magnitude better than this. Luckily, methods exist. Currently many sensors operate at SQL, a few operate ~10dB below. More on this in a few slides.

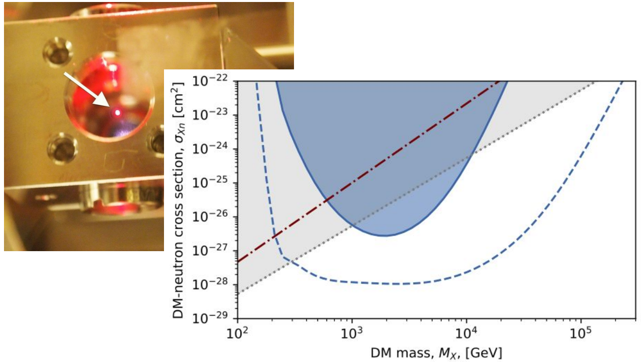
Some targets from SQL onward

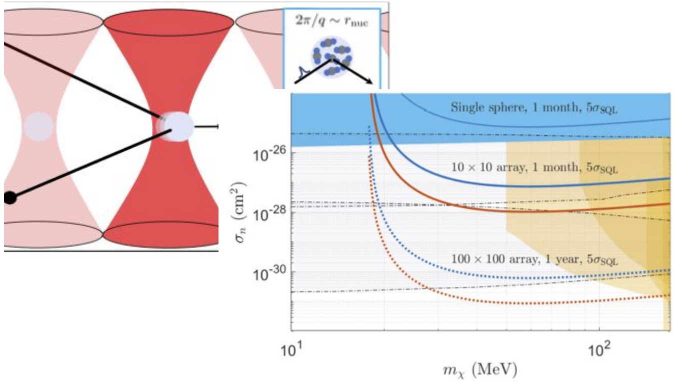


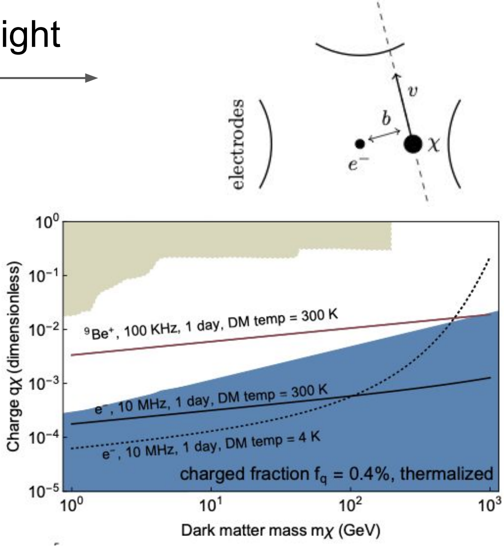
Impulses at various mass scales

detector mass

heavy → light







Composite DM coupled through long-range force [experiment]

~ug-scale levitated sphere

Monteiro, Afek, Carney, Krnjaic, Wang, Moore PRL 2020

Coherent elastic DM scattering at ~100 MeV mass [proposal]

~ng-scale levitated sphere

Afek, Carney, Moore PRL 2022

Rutherford scattering with milli-charged DM [proposal]

single trapped ion/electron

Carney, Haffner, Moore, Taylor PRL 2021

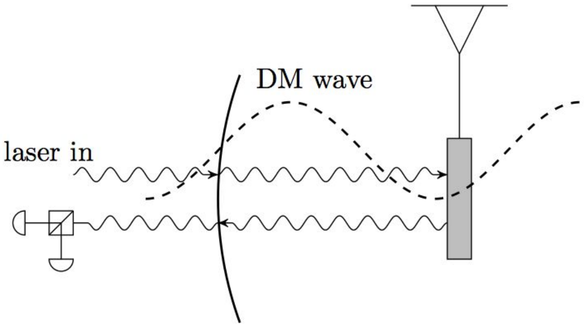
Ultralight DM detection

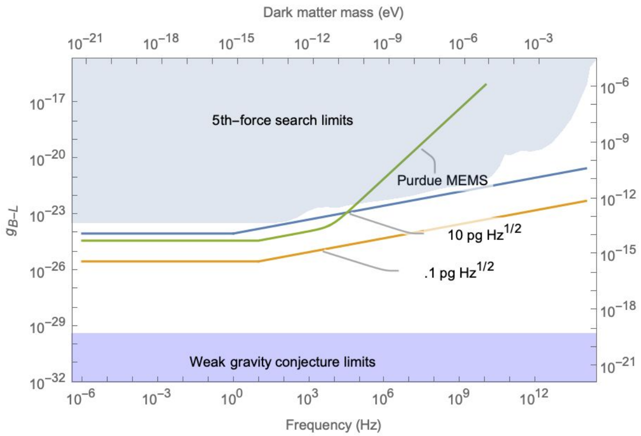
Example: DM m < 0.1 eV, coupled to B-L charge

Coherent, persistent, oscillating force on mechanical sensor → acceleration signal

$$\mathcal{L}_{int} = g_{B-L} \not{A} \bar{n} n \longrightarrow F = g_{B-L} N_n F_0 \sin(\omega_s t)$$

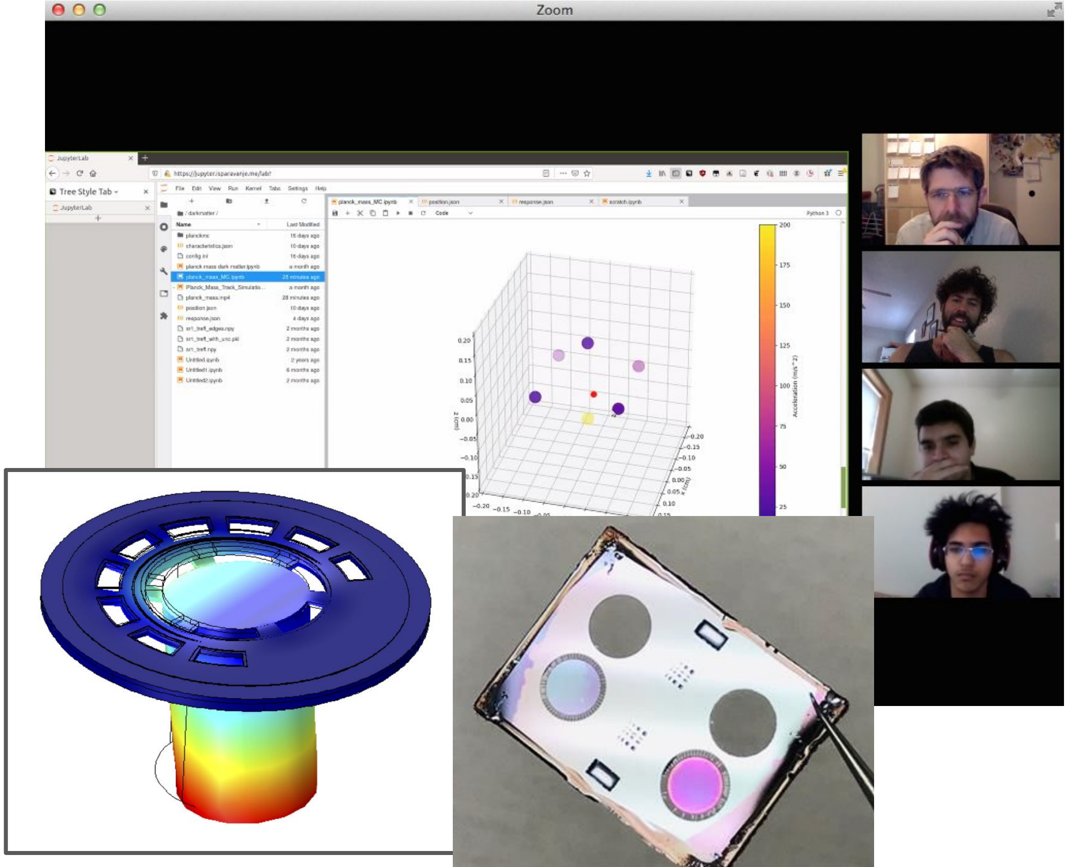
For comparison, LISA pathfinder had ~10^{-3} pg/rtHz. With N sensors, get full sqrt(N) noise reduction here.





1512.06165 Graham, Kaplan, Mardon, Rajendran, Terrano
1908.04797 Carney, Hook, Liu, Taylor, Zhao

Windchime



Rafael Lang @ Purdue

Collaboration currently involving Purdue, ORNL, Rice, FNAL, Maryland, Minnesota, NIST, LBL, …

Ultimate goal: gravitational detection w/ 10^6-10^9 sensors (!), << SQL

environmental isolation

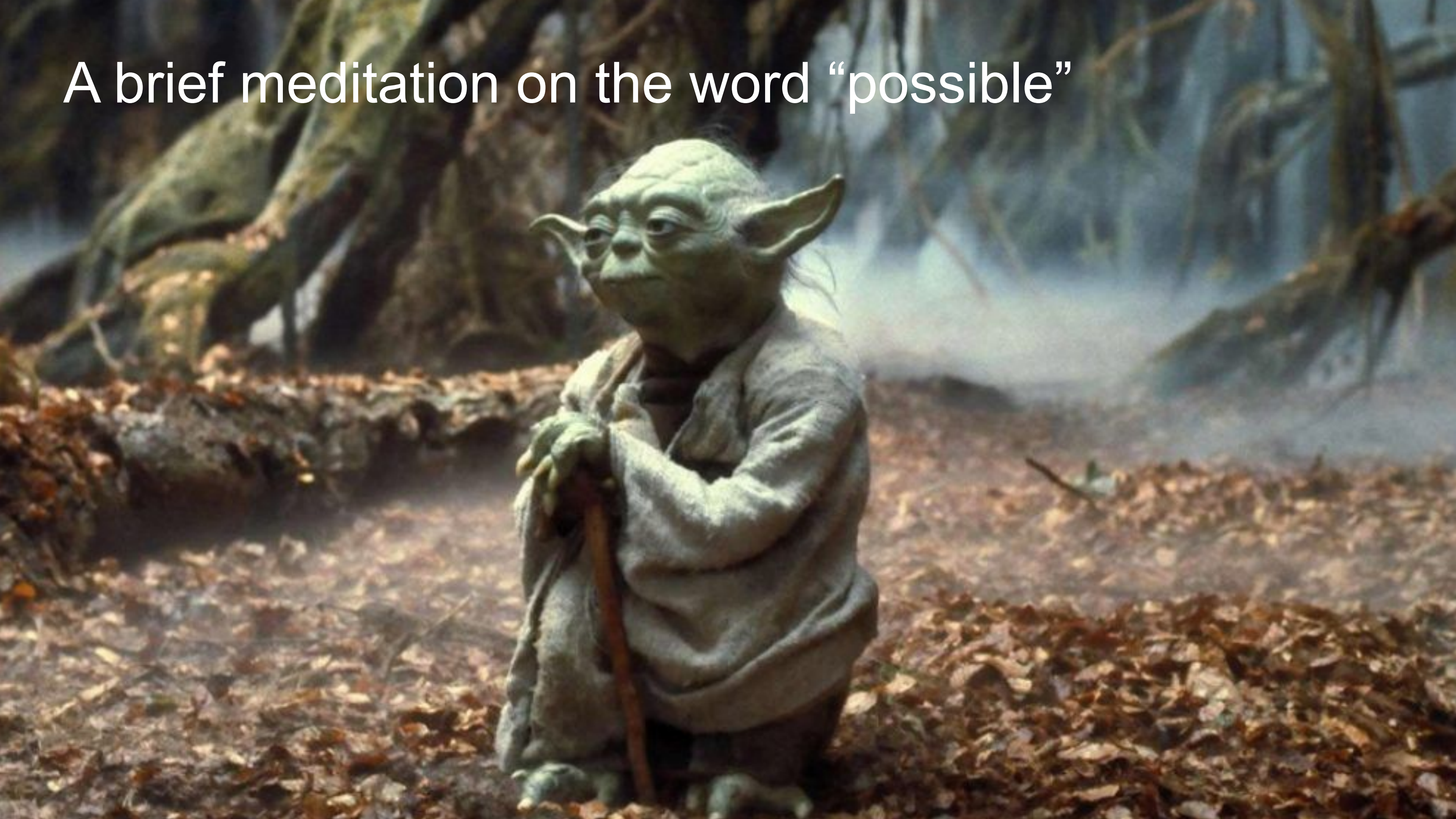
signal processing & computing

sensors

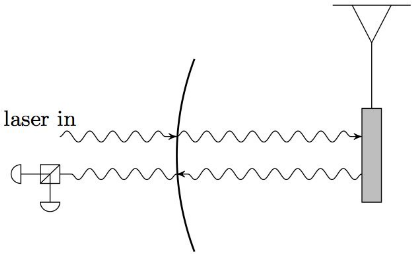
measurement noise

Snowmass: 2203.07242

A brief meditation on the word “possible”



Quantum mechanics and measurement

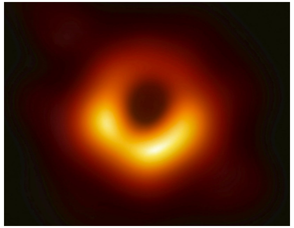


Gravitational DM detection appears to require measurements well beyond the “standard quantum limit”. Is this really possible?

It is not possible with any sensor we have now, and sufficiently scaling/tweaking any current sensor does not look plausible.

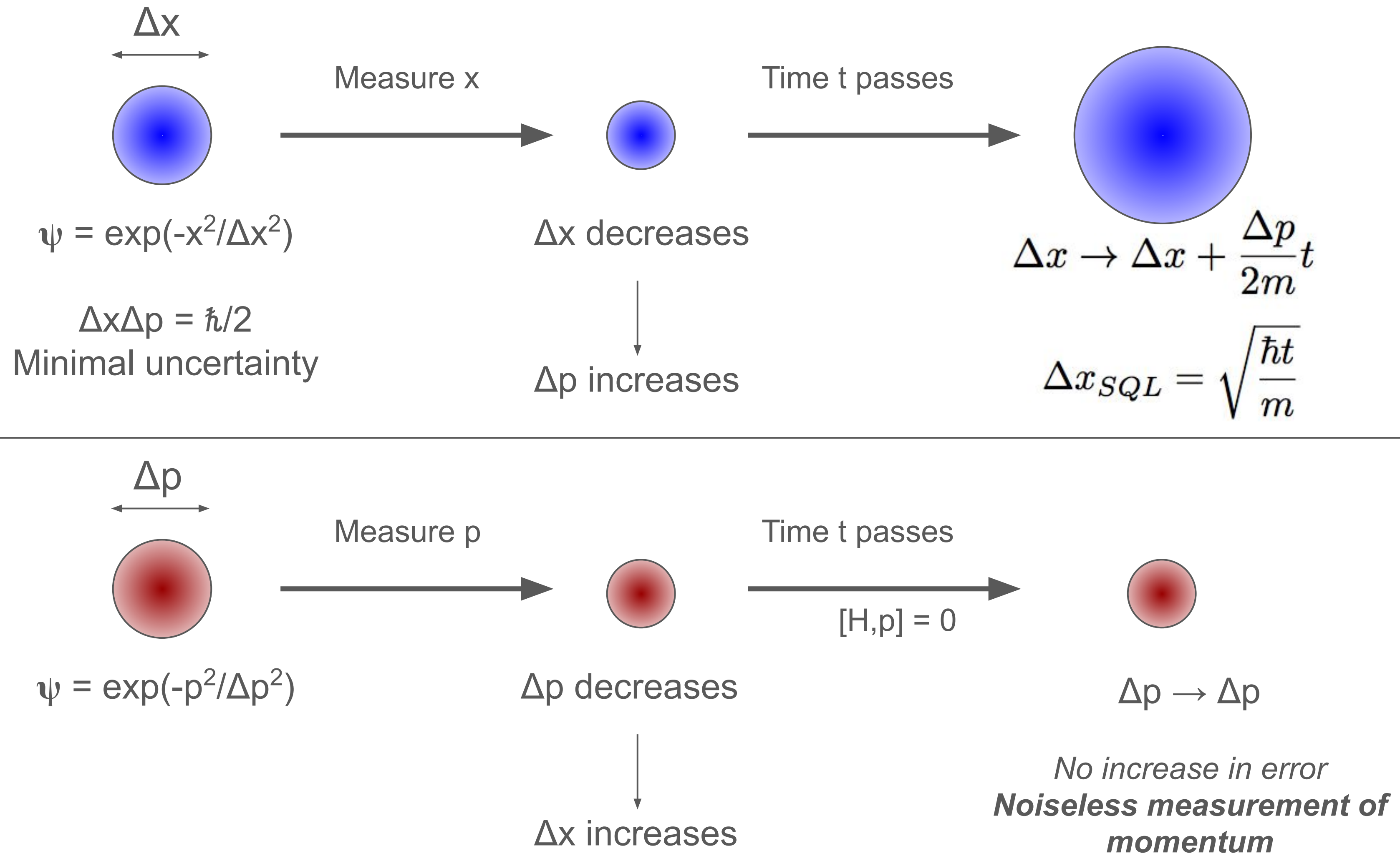
$\Delta x = 0 \rightarrow \Delta p = \infty$

But I think one should proceed without fear.

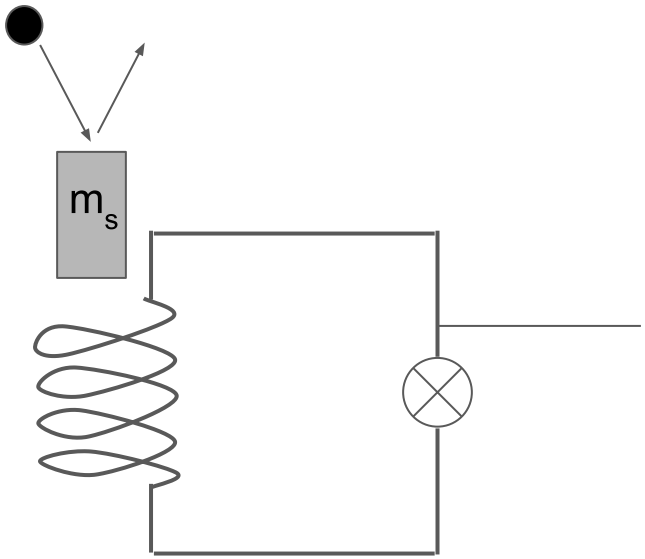


Quantum mechanics itself does not impose any limit to how precisely one can measure a system. The Caves argument given earlier can be evaded.

Ultimately, the only *fundamental* limit to what is possible in measurement will come from quantum gravity…


Δx
Measure x
Time t passes
$\psi = \exp(-x^2/\Delta x^2)$
Δx decreases
$\Delta x \to \Delta x + \frac{\Delta p}{2m} t$
$\Delta x \Delta p = \hbar/2$
Minimal uncertainty
Δp increases
$\Delta x_{SQL} = \sqrt{\frac{\hbar t}{m}}$
Δp
Measure p
Time t passes
$[H,p] = 0$
$\psi = \exp(-p^2/\Delta p^2)$
Δp decreases
$\Delta p \to \Delta p$
No increase in error
Noiseless measurement of momentum
Δx increases

Noiseless momentum measurements?

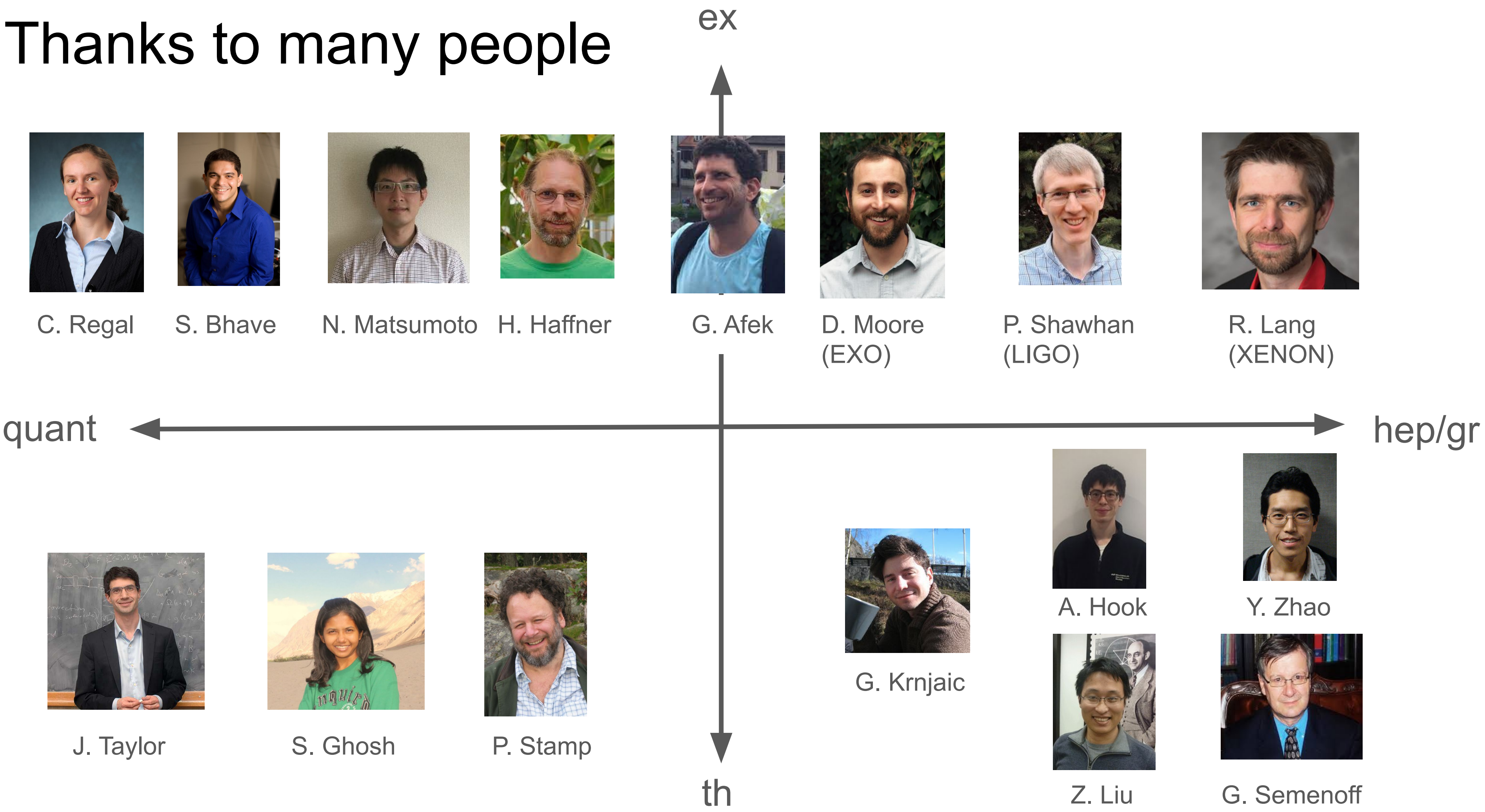


Moving magnet → EMF ~ v → read out via junction

$$\nabla \times \mathbf{E} = -\frac{\partial \mathbf{B}}{\partial t}$$

Unpublished work w/ S. Ghosh, B. Richman, J. Taylor
See also:
Caves, Thorne, Drever, Sandberg, Zimmerman RMP 1980
Ghosh, Carney, Shawhan, Taylor 1910.11892

Thanks to many people



Come to Berkeley!

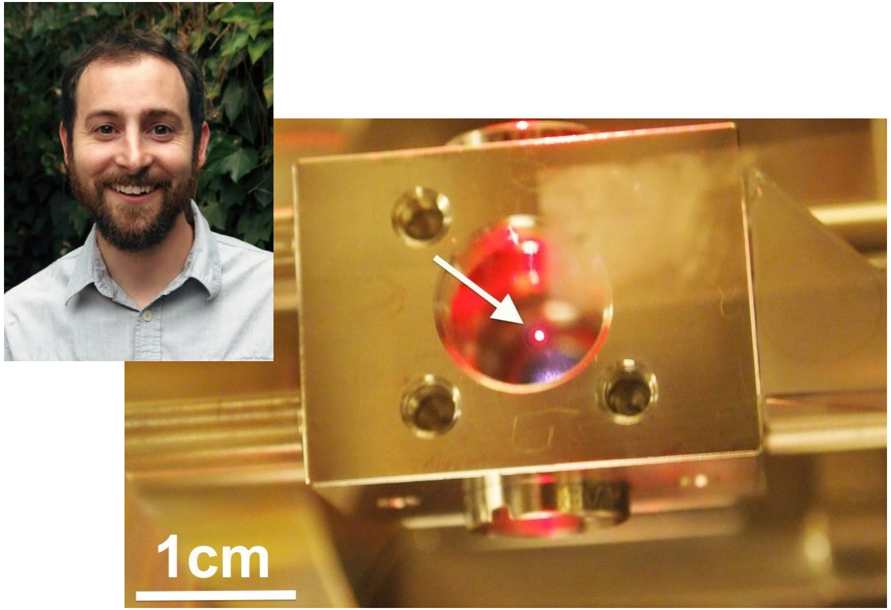
We are hiring a theory postdoc. Longer term hiring experimental people at different levels. Reach out for details: carney@lbl.gov

Outlook

- Gravitational direct detection of Planck-scale DM appears to be possible, but extremely technically challenging
- **Need better impulse sensing protocols -- theory and experiment**
- Between current experiments (~SQL level noise) and what we need (orders of magnitude below SQL), some clear physics targets have emerged: long-range coupled DM, light DM, neutrinos (another talk, sorry), ultralight DM, …
- At a high level, mechanical sensors are good whenever you want to look for a signal coherent across the size of the sensor. Any good ideas??

Backup slides

Yale experiments



Search for new Interactions in a Microsphere Precision Levitation Experiment (SIMPLE) @ D. Moore group

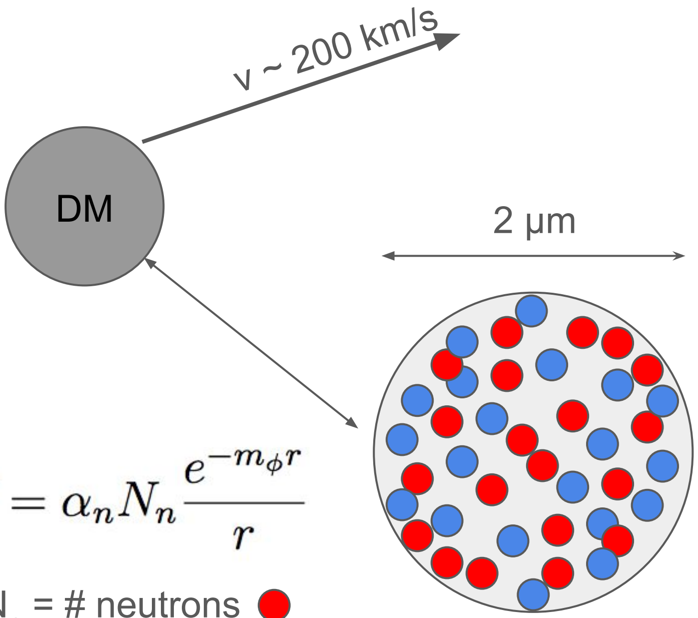
0.1-10 ng dielectric spheres

Optically levitated, stability ~ days

Continuous (biaxial) position monitoring at ~10x SQL level

Monitor this sphere for jumps in its motion. If it doesn't jump more than a few times then we can rule out DM models that would have caused jumps.

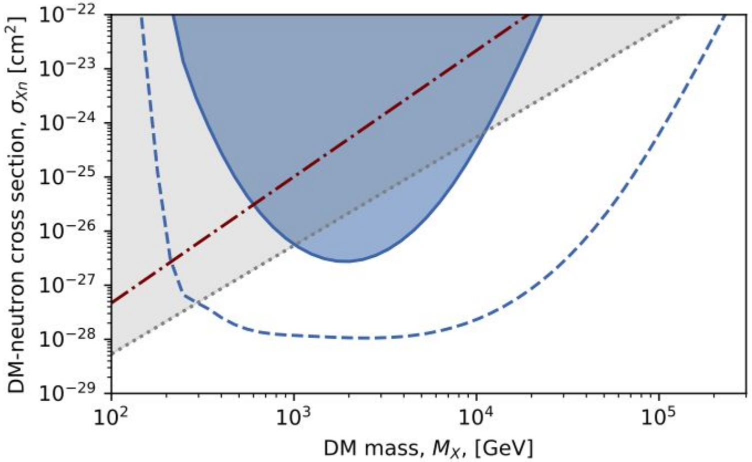
Composite DM with light mediator



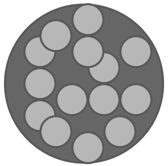
$$V = \alpha_n N_n \frac{e^{-m_\phi r}}{r}$$

N_n = # neutrons

sensor



One possible microscopic realization, “dark quark nuggets” coupled through B-L



Lin, Yu, Zurek 1111.0293
Krnjaic, Sigurdson 1406.1171

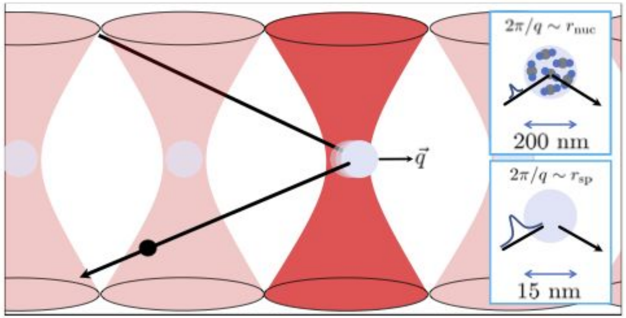
$$\alpha_n \to g_n g_d N_d$$

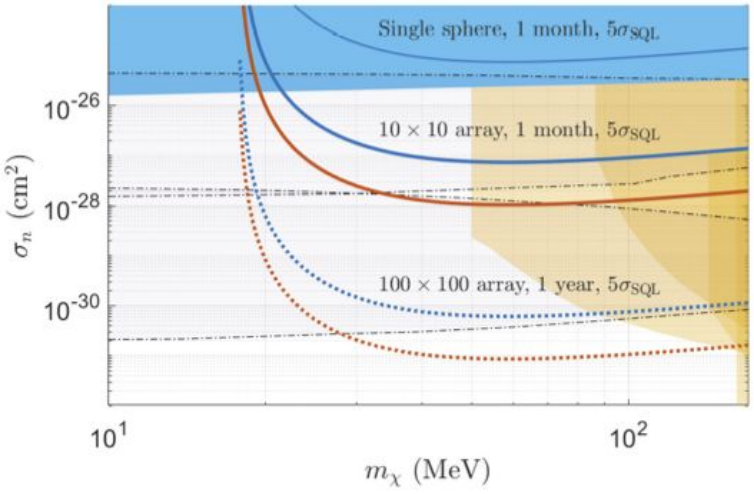
Nanospheres

What can you do with spheres ~ 1000 times smaller? (~10 nm)

Look for lighter dark matter!

In fact you can look for fundamental (non-composite) DM at this scale. It can scatter quantum-coherently off the sphere.





Proposal: Afek, Carney, Moore PRL 2022

Single ions… or electrons?

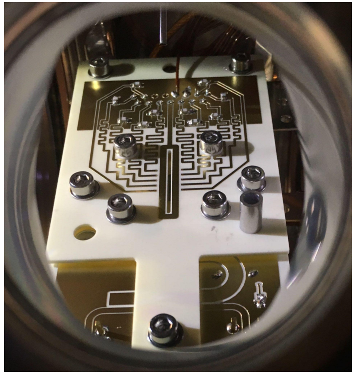
What's the ultimate limit of this idea, in terms of shrinking the sensor?

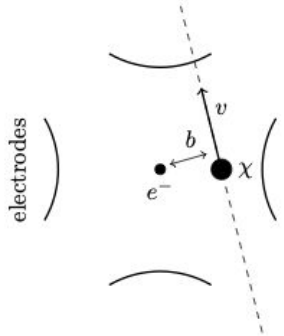
Using a single atom or electron!

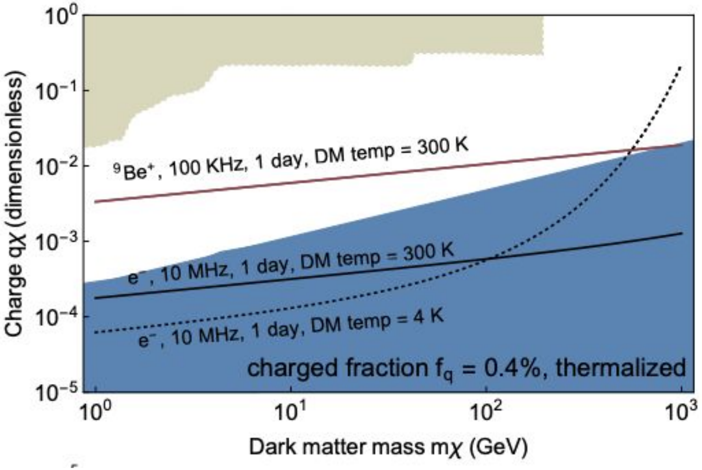
Can search for DM if it has tiny electric charge (~1/1000th the charge of electron).

Current detectors are totally insensitive to this regime

Carney, Haffner, Moore, Taylor PRL 2021







Matched filtering and SNR

Process the raw data via filter (cf. LIGO matching to waveform). For observable, use total impulse, filtered appropriately:

$$O(t_e) = \int f(t_e - t')F(t')dt'$$

$$f_{\mathrm{opt}}(\nu) = \frac{F_{\mathrm{sig}}(\nu)}{N(\nu)}$$

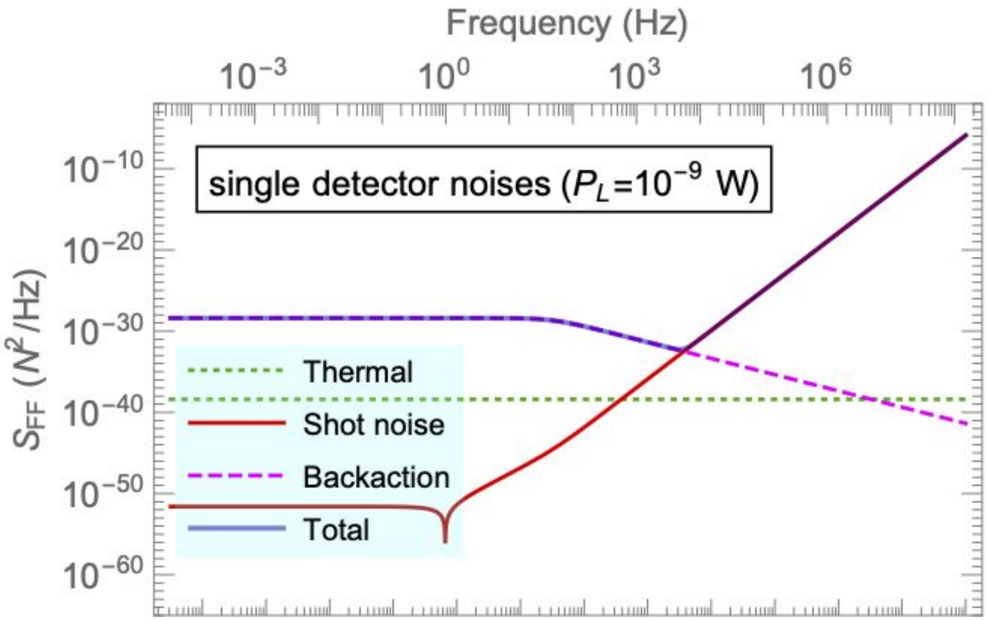
Known signal shape (e.g. F=1/r^2) and known noise power spectral density N, maximize SNR

$$\boxed{\mathrm{SNR}^2_{\mathrm{opt}} = \int_0^\infty \frac{|F_{\mathrm{sig}}(\nu)|^2}{N(\nu)} d\nu}$$

Limits on the noise

If we're looking for a signal with known shape (e.g. $F = Gm^2/r^2$), the best SNR possible is given by

$$\mathrm{SNR} = \sqrt{\int d\nu \frac{|F_s(\nu)|^2}{S_{FF}(\nu)}}$$



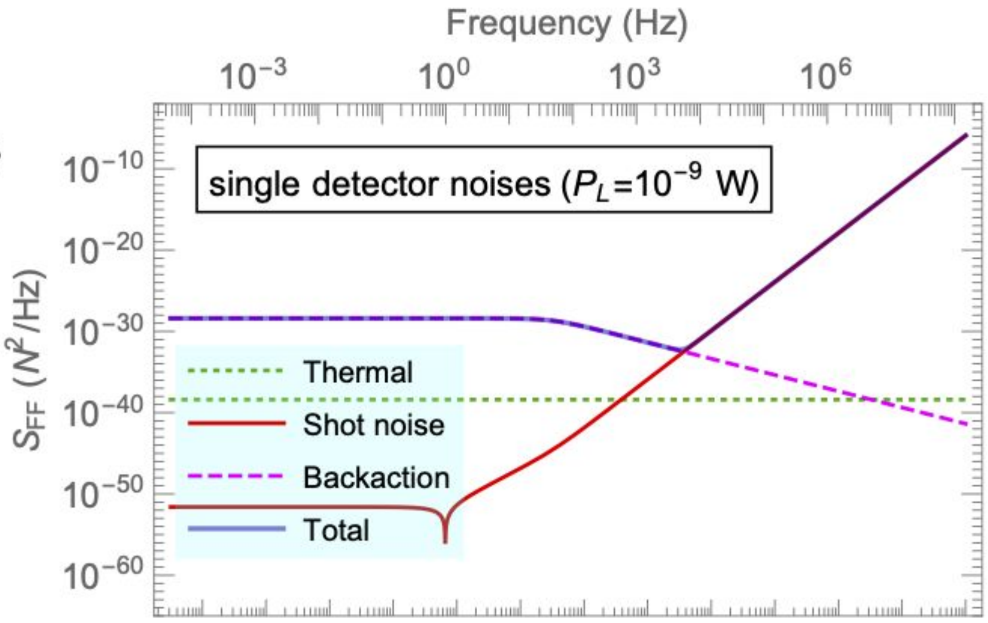
For impulses here:

$F_s(\nu)$~ flat up to $1/t_{\mathrm{flyby}}$ = v/b ~ 1 MHz - 1 GHz

$$Y_{\text{out}}(\nu) = A(\nu)Y_{\text{in}}(\nu) + B(\nu)X_{\text{in}}(\nu) + C(\nu)F_{\text{in}}(\nu)$$

$$F_{\text{E}}(\nu) = \frac{Y_{\text{out}}(\nu)}{C(\nu)}$$

$$S_{FF} \sim \langle F_{\text{E}}^2 \rangle \sim \frac{|A|^2}{|C|^2}\langle X_{\text{in}}^2 \rangle + \frac{|B|^2}{|C|^2}\langle Y_{\text{in}}^2 \rangle + \langle F_{\text{in}}^2 \rangle$$



$$S_{FF}^{SN} = \frac{x_0^2}{G^2\kappa}\frac{1}{|\chi_c\chi_m|^2}$$

$$S_{FF}^{BA} = \frac{G^2\kappa}{x_0^2}|\chi_c|^2$$

$$S_{FF}^{T} = 4\gamma m_s kT.$$